HERSCHEL SPACE OBSERVATORY

Announcement of Opportunity for Open Time Programmes

Herschel Observers' Manual

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Preface

The Herschel Space Observatory is an ESA cornerstone mission that was be launched on 14 May 2009, alongside the Plank cosmic microwave background mission. Originally known as FIRST (Far InfraRed Submillimetre Telescope) its name was officially changed in the year 2000 in recognition of the 200th anniversary of the discovery of infrared radiation by William Herschel in 1800. Herschel covers the range from 55 to 672 microns (530-5000GHz) - a region that is effectively totally closed to ground-based astronomy - using a suite of three state-of-the-art instruments called PACS, SPIRE and HIFI.

Herschel is an observatory mission: that is, its time is distributed among the community instead of being used for a large-scale survey. It is also a consumables-limited mission - its useful life depends on the lifetime of the helium in the dewar that is used to cool the instruments and is expected to be in the range from 3.5 to 4 years from launch. As an observatory mission its success thus depends on the quality of the science that the community carries out with it and how effectively the helium in its dewar is converted into science. The "helium into science" ratio will be the principal deciding factor in allocating time with the Herschel Space Observatory.

Many aspects of the Herschel Space Observatory are revolutionary. It is, thanks to its innovative design, the largest dedicated infrared telescope ever to be launched into space by a considerable margin. For the astronomer this converts into high sensitivity and a spatial resolution a factor of 6 better than any previous far-infrared telescope launched into space, making Herschel a pathfinder mission in the far-IR. In fact, Herschel is limited in sensitivity mainly by the confusion from the background of faint, unresolved sources. This makes Herschel a revolution for astronomy in a range of the far-IR that has hardly been exploited so far. Herschel observations will have a huge impact on astronomy and on our understanding of the universe.

This manual describes the observatory aspects of the mission: the spacecraft and its performance; the mission; the space environment in which the Herschel Space Observatory is operating (very different from previous missions such as IRAS, ISO and the HST); and use of Herschel - from how an observing proposal is received and treated, through to final archiving of the data. The aim is to give an overview of Herschel to the user, describing everything that a potential observer needs to know at a superficial level; where deeper knowledge is required afterwards, the observer should go to the specific documentation for each system or sub-system (e.g. the individual instrument manuals, the Data Processing user manual, etc.) The aim is that simply by reading this manual, or by using it for reference, someone who is planning to request time with Herschel has enough information to decide whether or not to proceed and to have a clear idea how to start.

When this manual was first written for the Guaranteed Time Key Programme Call back in November 2006, the launch of Herschel was still 30 months away and knowledge of how the spacecraft and instruments would behave in space was theoretical. Similarly, some important elements of the Science Ground Segment were still in development. At the time of this revision we are 22 months into the mission and have characterised most aspects of Herschel's performance and operation thoroughly. As a result, this manual has undergone a further deep revision and its contents have been updated to reflect what is now a quite mature operational reality.

Chapter 1. Mission phases

Herschel flight operations are divided into a series of phases from the moment of launch. Broadly, these are check-out, routine operations and post-operations, each with their individual sub-phases. In theory, each mission phase should have an exact start and end point but, in reality, the requirements of operations and the differing needs of the three instruments have made the different mission phases blend slowly into each other, with slow transitions and no clear start and end point. Similarly, the HIFI anomaly meant that HIFI was delayed by about 6 months with respect to PACS and SPIRE in entering routine operations and that it had to return over its tracks for a time and re-conduct check-out activities that had already been completed months earlier.

Overall, Herschel operations have run very smoothly, largely due to the success of the long and intensive pre-flight test campaign process, first at ESTEC and later at Kourou. As a result, some activities could be advanced considerably over the anticipated pre-launch schedule. Overall, remarkably few check-out activities failed for such a complex mission, thus requiring much less re-planning of in-flight tests than might otherwise have been expected.

1.1. Completed mission phases

1.1.1. Early mission history

Roll-out (shown in Figure 1.1), prior to launch, was conducted early on the morning of 13 May 2009 and, after a flawless countdown, launch occurred at 13:12UT on 14 May. Although there had been storms and heavy rain as the guests were being transported to the VIP area, the clouds disolved and launch conditions were perfect. Figure 1.2 shows the Ariane 5 blasting off with the Herschel and Planck on board. The critical early milestones of fairing release, Herschel separation (at 13:37:55UT) from the Scylda and signal acquisition (at 13:49UT) were all passed successfully. The frst command was executed by Herschel 58 min after launch at 14:10UT and Herschel began its cruise to Lagrange.

Figure 1.1. Roll-out of the launcher for the Herschel-Planck mission on 13 May 2009.

Figure 1.2. Launch of the Herschel-Planck mission on an Ariane 5-ECA at 13:12UT on 14 May 2009.

Figure 1.2 shows Herschel, Planck and the Sylda approximately 26 hours after launch, when they were already 226 000km from Earth (approximately 0.5 Lunar Distances). It is already obvious from this image how Herschel and Planck were released into slightly different transfer orbits due their differing injection requirements for L2 orbit. A single injection manoeuvre was made 26h after launch. This injection manoeuvre started at 15:16:25UT and lasted 22.5 minutes, giving a Delta-V of 8.7m/s. This injection manouevre placed Herschel in what was effectively its final orbit and was

so successful that no further corrections were required apart from the tiny, regular station-keeping burns of typically 10-20cm/s with the thrusters, made every 4-6 weeks to maintain the orbit around L2.

HERSCHEL, PLANCK and SYLDA
226,000 km from Earth on 2009 May 15
Faulkes Telescope South   (R. Miles)
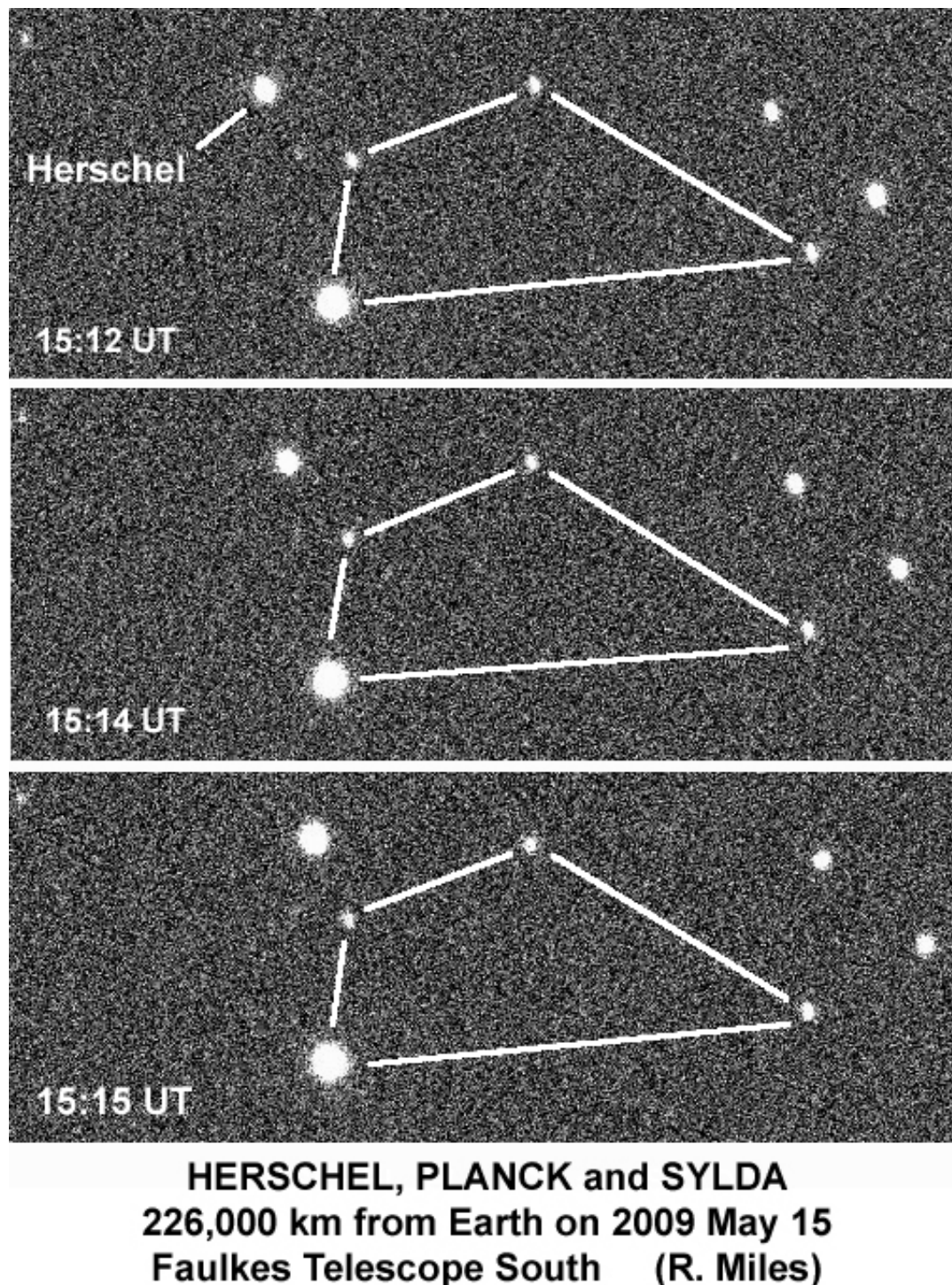
Figure 1.3. A sequence of images taken by British amateur astronomer Richard Miles using the 2-m Fawkes South Telescope in Australia of Herschel (identified), Planck and the Sylda 26 hours after launch, at approximately half the distance to the Moon.

After launch a Low Earth Orbit Phase (LEOP) started, an initial phase of check-out started with the telescope closed and the instruments switched off while the operation of the spacecraft sub-systems was checked. At this time the satellite was placed in a Sun-orientation allowing the spacecraft to cool in the shadow of the sunshade and to outgas. However, given the danger of volatiles from the

satellite condensing on the telescope, the mirror itself was heated during the initial cooldown phase to avoid it acting as a cold-trap for the outgassed volatiles. This period involved checking basic properties of the satellite (centre of mass, moments of inertia) and proper functioning of basic spacecraft sub-systems (Radio Frequency (RF), thermal control, power sub-system, data handling, attitude and orbit control, thrusters, Solid State Recorders (SSR), etc.), at least to the extent that these sub-systems were required for spacecraft operations.

While spacecraft check-out was underway, initial activities to check-out the instruments could commence. The first instrument to be switched on to start payload operations was SPIRE on Day 6 after launch, with PACS and HIFI switch-on on Day 11. Initial switch-on simply consisted of checking that the measured voltages were in the expected range from the telemetry. This led into a long and extremely detailed set of tests and checks of the functionality of each instrument that was the Commissioning Phase.

1.1.2. Commissioning Phase

Once Herschel was successfully launched and injected into the transfer trajectory towards the operational orbit, the spacecraft and instrument commissioning phase started. This consisted of a series of 298 individual tests and activities to check-out all aspects of instrument and spacecraft functionality. A highlight of Commissioning was the opening of the telescope cryo-cover. This cover protected the cryostat from condensation of outgassed volatiles. The cryo-cover was opened at 10:53UT on Sunday 14 June (12:53 local time at Darmstadt). This involved firing explosive bolts to free the cover after which a spring pulled it into an upright, totally open position after a series of oscillations during which a spring steadied the heavy cryocover into a fully open position. The oscillations of the cover caused the gyros to activate to stabilise the spacecraft pointing because the cover was heavy enough to make the whole satellite move slightly in reaction (shown in Figure 1.4).

Figure 1.4. The telemetry received at MOC showing the oscillation in gyro response as the heavy cryocover swung open and oscillated, causing the entire satellite to wobble slightly until it had reached a stable open position. This gyro signal was the first evidence that the cryocover opening had been carried out successfully.

With the successful cryocover opening and the encouraging progress of commissioning activities, an opportunity was seen to take some early images to make a blind test of the telescope focus and image quality in advance of formal First Light. A series of PACS exposures were defined with a range of bias settings, scanning through the most likely range of values in a test that was termed "Sneak Preview". The resulting images are shown in Figure 1.5, which showed that the telescope focus and alignment were excellent and that the optimum parameters for imaging were close to the best guess values estimated by PACS prior to cryocover opening.
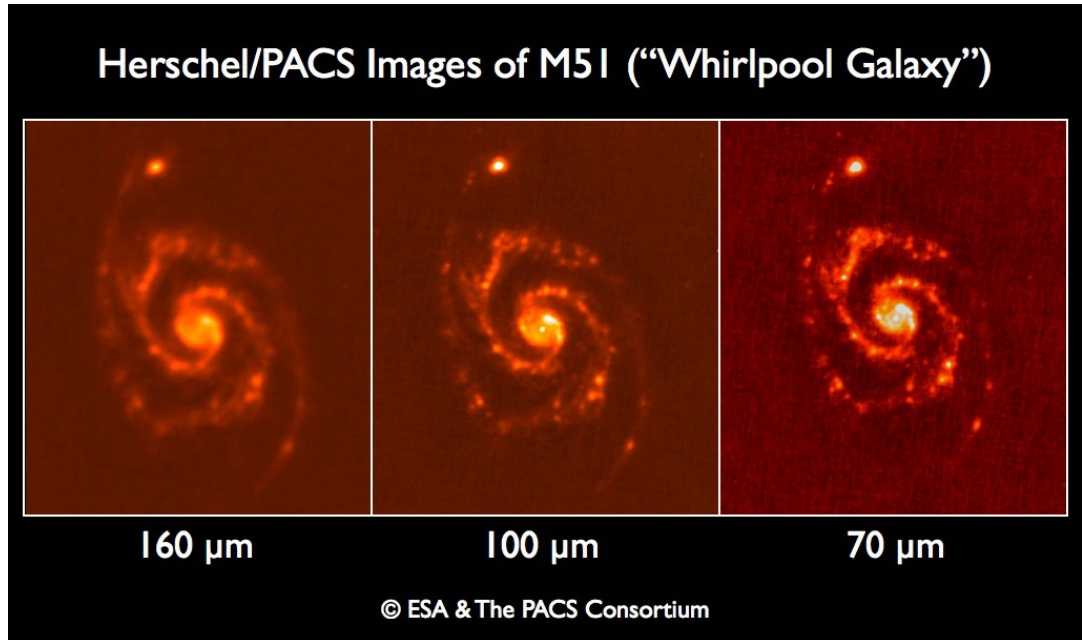
Figure 1.5. The Sneak Preview images of M51 in the three PACS bands, taken blind after cryocover opening.

M51 was selected for several reasons, especially the fact that, apart from being available at the right time and being a large, bright target, the galaxy is a classic infrared target with a lot of structure, so considerable prior imaging existed at similar wavelengths (Figure 1.6) that could be used to check the image quality and ensure that it met expectations. Once Sneak Preview had shown that the image quality met all expectations, a more ambitious series of formal first light observations were scheduled for each of the instruments to demonstrate their capabilities (Figure 1.7, Figure 1.8).
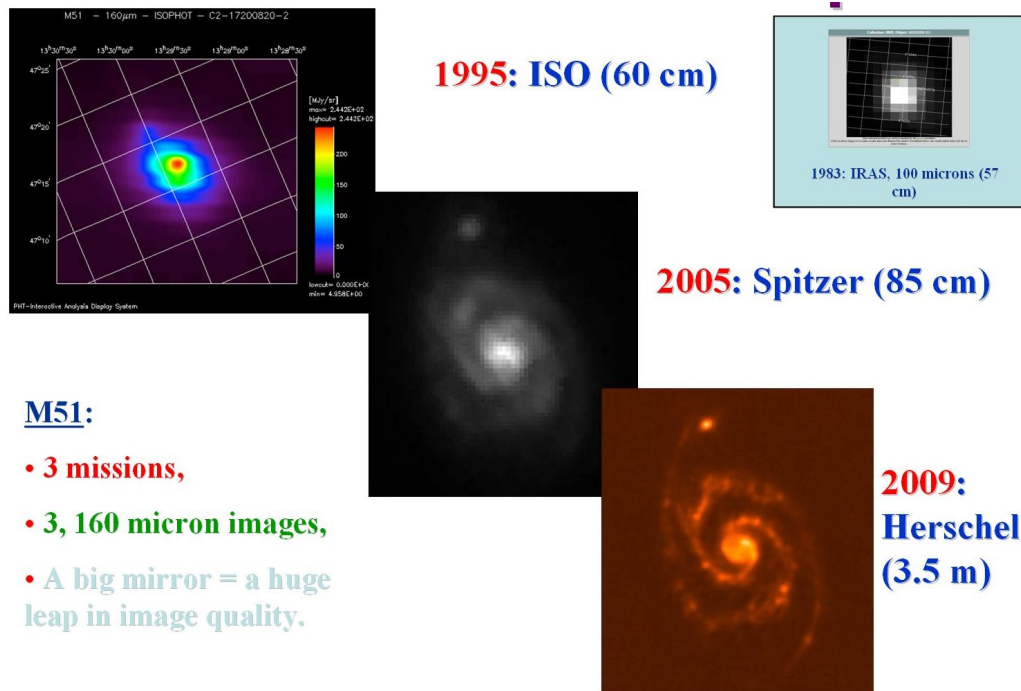
Figure 1.6. A comparison of images of M51 at 160 microns for ISO, Spitzer and the Herschel sneak pre-view image, showing the improved resolution and sensitivity from Herschel's larger mirror. No comparable image exists from IRAS, so the IRAS 100 micron image is shown for comparison.
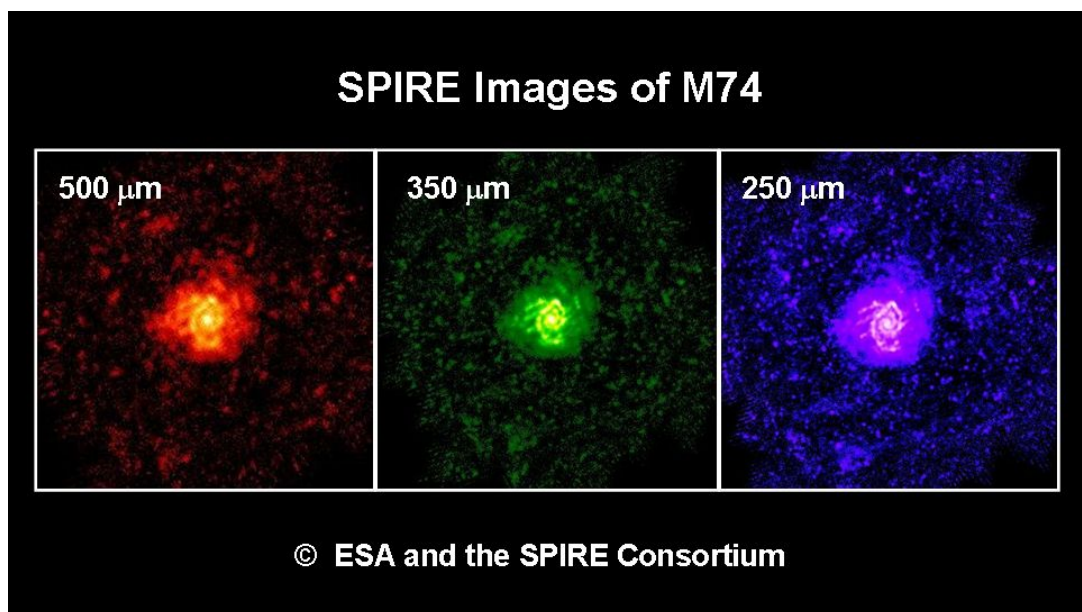
Figure 1.7. The SPIRE First Light images of M74 in 250, 350 and 500 microns, obtained on 2009 July 9.
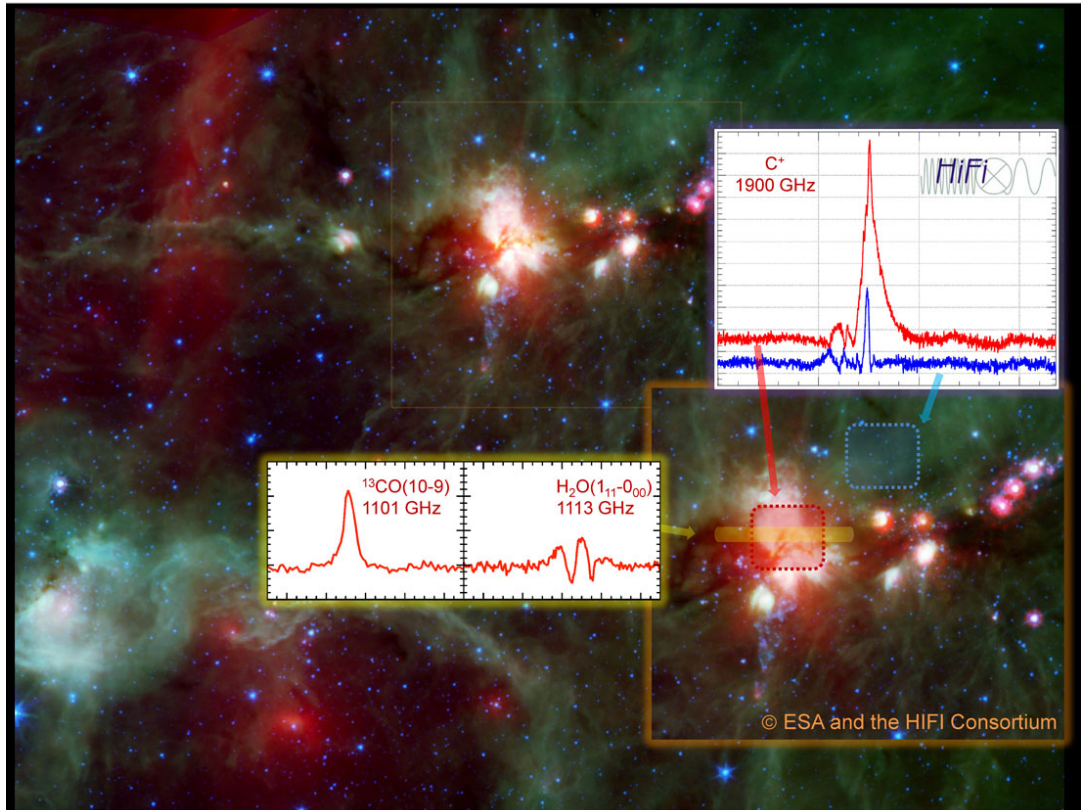
Figure 1.8. The HIFI First Light spectra of DR21, obtained on 2009 July 9.

1.1.3. Performance Verification (PV) Phase

PV phase was designed to obtain in-flight characterisation of all instruments e.g. in terms of stability, sensitivity, resolution, timing and other calibration parameters. It included the validation of the instrument observing modes and the calibration and data processing of the resulting data. To achieve this, a schedule of astronomical observations and internal calibrations, defined and iterated pre-launch covering a nominal period of 2 months were be executed using normal observatory procedures. This schedule was be based upon an agreed in-orbit calibration plan generated jointly by the ICCs and the HSC. The plan contained a description of all planned calibration activities and associated calibration sources (internal and astronomical) required to characterise fully each instrument.

Each instrument received blocks of time, normally of two days each, to carry out its activities according to the agreed PV plan, giving each instrument "two days on and four days off", allowing data to be processed and new observations prepared. Weekly meetings then examined the progress of the planned tests, adjusting the plan to allow extra time for failed tests to be repeated, where necessary, or for extra tests to be included. PV Phase started 64 days after launch - in line with pre-launch plans - and, by 120 days after launch had delivered the first fully calibrated and usable observing modes for science scheduling.

PV Phase blended progressively into the Science Demonstration Phase of Routine Operations, without a formal end, with days assigned to PV activities becoming increasingly infrequent. Apart from HIFI recovery activities to check that the instrument was functioning correctly on its back-up chain after the incident in July 2009 that stopped observations, which ran in four dedicated blocks between 22 January 2010 and 17 March 2010, the last PV day included as such in the observing schedule was 8 December 2009. After this, remaining PV activities to test and validate remaining observing modes have been absorbed into routine calibration activities for each instrument, as have the activities designed to test and validate the new, second-generation observing modes that have been offered since.

1.1.4. Science Demonstration Phase (SDP)

As noted above, the PV Phase blended progressively into the "Science Demonstration Phase", in which each approved Herschel science programme had the opportunity to nominate a part of its observations -- typically 5-10% -- to be carried out early. The aim was to carry out observations and observing programmes that would test the capabilities of Herschel in detail, frequently with difficult and challenging observations. This allowed astronomers to test their observing strategy, compare data quality with expectations and fine-tune their observing programmes, as well as allowing a global overview to be obtained of the performance of the three instruments.

Observations were carried out on a shared-risk basis: astronomers could opt to forego their proprietary rights on data and allow them to be made public at the opening of the Herschel science archive and, in return, would get the time used re-imbursed in their programmes by Herschel; alternatively, they could maintain the data proprietary for one year from execution and have the data counted as part of their Routine Science programme. So, in return for assuming part of the risk of testing observing modes early in the mission, astronomers had the chance to obtain early publication of Herschel data, selecting their most critical observations for rapid execution.

Science Demonstration Phase demonstrated that the main mapping modes that are the workhorses of Herschel were essentially ready to go, particularly in its early mission phases, although some extremely useful input was obtained for observing strategies, leading to the recommendations on how to obtain the best sensitivity in observations. Similarly, it gave a lot of valuable information on how best to define spectroscopic observing modes. Early results from Science Demonstration Phase were presented at the Herschel "Initial Results" Workshop at Madrid in December 2009 at which every approved Herschel project had some observational data to present.

The prime period for SDP was defined to be from 15 September to 15 December 2009. Those SDP observations that could not be completed by 15 December because of target visibility concerns, or because the required observing mode had not been released, were given the highest priority for telescope scheduling over the period up to 30 April 2010, when the few remaining SDP observations that had not been scheduled reverted to being treated as Routine Science observations.

1.1.5. HIFI Priority Science Programme (PSP)

After the HIFI anomaly, it was decided to define a variant of SDP for HIFI observers to be executed as rapidly after HIFI recovery as time permitted. This was the HIFI Priority Science Programme, or PSP. Two blocks of telescope time in March and April were reserved for intensive HIFI observing campaigns. Observations were divided into PSP1 (highest priority) and PSP2 (second priority) to fill these blocks of time efficiently, allowing a substantial part of the Herschel's HIFI observations to be carried out at the earliest possible date, after which HIFI would enter the standard Mission Planning cycle with a set number of days assigned each 4 weeks, with top priority going to scheduling remaining PSP observations. A special HIFI initial results workshop was arranged in Leiden in April 2010 to present a first look at PSP data and checkpoint for observing strategies.

1.2. Current and future mission phases

1.2.1. Routine operations (Routine Phase)

As with previous mission phases, there is no clear transition between SDP and routine operations. As each project received its SDP data, if no significant problems were revealed, the Principal Investigator (PI) was invited to have a release telecon with the Project Scientist and HSC staff to discuss the data and any problems or issues that had arisen. If no serious issues were identified, the PI was invited to release all, or part of the observations in his or her programme for scheduling, in which case, the observations would be made available to the HSC Mission Planners. The first routine observations were observed on 18 October 2009 and, by December, the immense majority of scheduled observations came from released routine programmes. Over the course of the mission Herschel will produce hundreds, or thousands of spectacular images like these Figure 1.9, Figure 1.10.

Figure 1.9. M31's once and future stars. A combined Herschel and XMM image of M31 showing dusty start-forming regions (Herschel) and the point-sources that represent highly evolved stars (XMM). The Herschel data were taken at 250 microns with SPIRE between 17 and 21 December 2010. In the XMM RGB image, red sources are low-mass x-ray binaries, while the blue sources are compact binaries with a neutron star or black hole secondary.
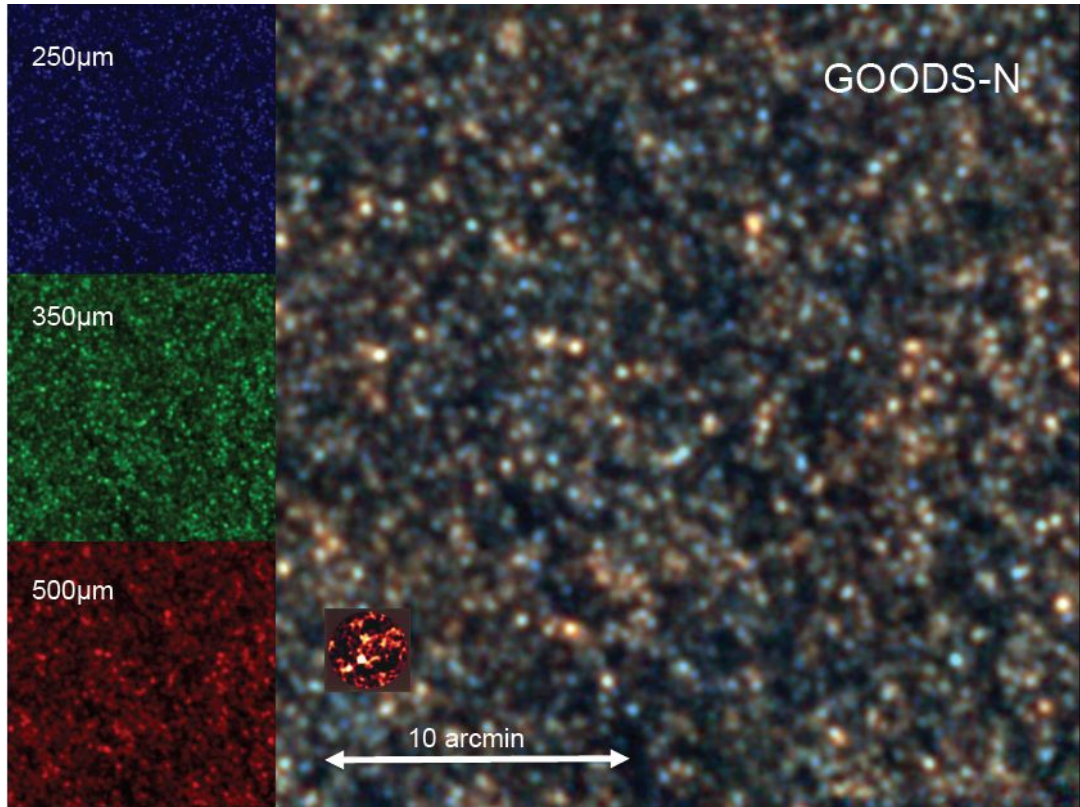
Figure 1.10. Galaxies spread like grains of sand on a beach. Every source in this GOODS-N field, which is about the size of the Full Moon, is a distant galaxy. The insert to the left shows the indivdual frames in each of the SPIRE bands while the main image combines them as an RGB. The colour of the galaxy gives an indication of its red shift and, hence, distance: the reddest galaxies are the most distant and may be as much as 12 000 million light years away; blue objects are relatively nearby and may be as close as 8000 million light years.

Herschel will carry out routine science operations phase for a minimum of 3 years. Early on, mainly Guaranteed Time and "Key Project" observing programmes have received priority. Key Projects were performed early in the mission to permit follow-up and to give the Guaranteed Time holders at the HSC the opportunity to obtain real data to work with, in preparation for supplying community support to the open time observers with the benefit of a thorough knowledge of the entire observing chain from proposal submission to access and reduction of data. Almost all Key Programme observations are expected to be completed by May 2011, at which point OT1 observing programmes will start to be heavily scheduled, although some OT1 observations have been scheduled since December 2010 where they help to improve efficiency by filling inconvenient gaps in the telescope schedule.

All observers can track the state of their proposals from the (password protected) proposal handling pages of the HSC Web page and are notified when the resulting data has been passed through the Quality Control process; this may take from 2-3 weeks to complete, although data is available for retrieval from the HSC usually within 48 hours of the observations being executed. Observers can also check both what observations are scheduled for observation and have been delivered to MOC (http://herschel.esac.esa.int/observing/ScheduleReport.html) and the observing log (http://herschel.esac.esa.int/observing/LogReport.html) from the HSC. Observations marked "Failed" are automatically cloned and released for re-scheduling by HSC.

1.2.2. Boil off

There is considerable uncertainty about when boil off will happen. This is unlikely to be reduced much in the future. Current best estimates place the most likely date for helium exhaustion around the end of 2012, but with an uncertainty of several months. It is possible that in the last few weeks of the mission it may be difficult or impossible to re-cycle the instrument coolers after the helium

level has dropped below a certain point and that, as a result, only spectroscopy will be schedulable, but this is still uncertain at this time as we do not know what the helium behaviour will be at very low levels.

Once boil off has occurred Herschel's instruments will no longer operate. There is still considerable discussion about what will happen to Herschel finally post-helium. MOC will continue to operate Herschel for a time even after boil off has occurred for routine spacecraft housekeeping operations. As the radiation monitors will continue to function indefinitely there is a suggestion that Herschel could continue to function after MOC operations have ended as a space weather station. Alternatively, it could be allowed to escape into solar orbit, or even crash into the Moon.

Table 1.1. Herschel mission key dates. Only approximate dates can be assigned to the different mission phases as there is inevitably a progressive transition between mission phases rather than a sharp one; in extreme cases there may be activities from three different mission phases progressing simultaneously and, in some cases, the start and end of a phase is a matter of definition and different dates could be given to the ones that appear here. In particular, HIFI recovery activities meant that CoP and PV days were scheduled months after the nominal end of these phases. Similarly, as reflected by this table, occasional PV days were being scheduled for PACS and SPIRE long after even routine observations had started.

Mission phase	Approximate Start	Approximate End
Launch	L=14 May 2009	
Early Orbit Phase	L	24 May 2009 (L+10 days)
Commissioning Phase	L	July 19th (L+66 days)
Performance Verification Phase	17 July 2009 (L+64 days)	25 November 2009 (L+195 days)
Science Demonstration Phase	11 September 2009 (L+120 days)	30 April 2010 (L+352 days)
Herschel Routine Phase	18 October 2009 (L+157 days)	L+36 months (current best guess); Boil-off = B
Run-down phase (3 months)	B	B+3 months
Mission consolidation phase (6 months)	B+3 months	B+9 months
Active archive phase (48 months)	B+9 months	B+57 months
Archive consolidation phase (6 months)	B+57 months	B+63 months (End of Herschel mission)
Historical archive phase (indefinite)	B+63 months	(TBD) End of all Herschel activity

1.2.3. Post-Operations Phase

The Herschel post-operations phase will consist of the rundown monitoring phase (starting at the moment of helium boil-off), mission consolidation phase, active archive phase, and the archive consolidation phase (at which point the transfer to the subsequent historical archive phase takes place), which is the final formal phase of the mission. Herschel is currently fully funded for 5 years of post-operations.

The goal of this phase is, within the constraints of time and available resources, to maximise the scientific return from the Herschel mission by facilitating continuing widespread effective and extensive exploitation of the Herschel data. This will continue after the conclusion of this phase (i.e. in the historical archive phase). Documentation will be extensively revised during this phase to provide a legacy and the Herschel Interactive Processing Environment (HIPE) will continue to be updated and refined in line with the state of knowledge of instrument behaviour and calibration.

The ultimate legacy of Herschel will be the historical archive, plus the sum of all the knowledge, both scientific and technical, derived from implementing and operating Herschel.

1.2.4. Archive Phase

The historical archive phase is outside the funded Herschel mission. This phase commences after the

end of the post-operations phase.

The historical archive will provide access to all Herschel observations and derived products. The products will all be derived in the archive consolidation phase during the post-operations phase in a consistent manner and to consistent standards using the best knowledge of Herschel instrument calibration and data processing. In addition, the software, documentation - manuals, etc.- and tools will be available from the historical archive.

Chapter 2. The Observatory

This section summarises the main characteristics of the Herschel spacecraft, its orbit, pointing performance and observable sky regions.

2.1. Spacecraft overview

The Herschel spacecraft has a modular design, comprising the *Extended Payload Module (EPLM)* and the *Service Module (SVM)*. The EPLM consists of the PLM "proper" with a superfluid helium cryostat - based on the proven ISO technology - housing the Herschel optical bench (HOB) with the instrument focal plane units (FPUs), and supporting the telescope, the sunshield/shade, and payload associated equipment. The SVM houses "warm" payload electronics and provides the necessary "infrastructure" for the satellite such as power, attitude and orbit control, the onboard data handling and command execution, communications, and safety. Figure 2.1 shows the main components of the Herschel S/C. Table 2.1 presents the Herschel Spacecraft key characteristics.
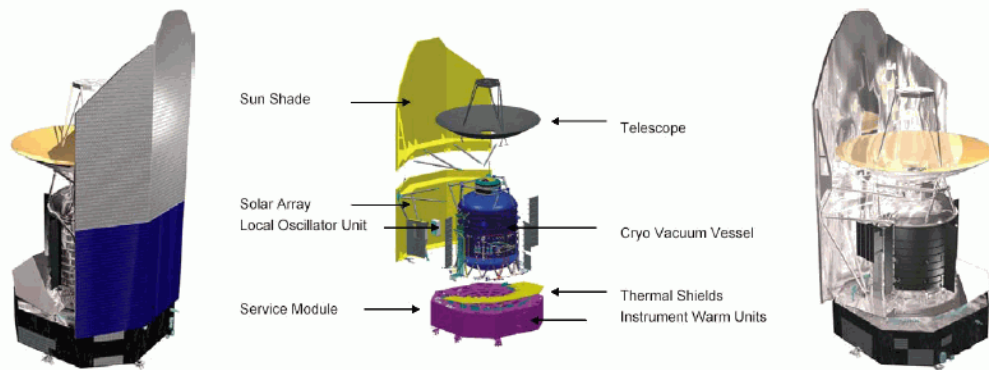
Figure 2.1. The Herschel spacecraft has a modular design. On the left, facing the "warm" side and on the right, facing the "cold" side of the spacecraft, the middle image names the major components.

Table 2.1. Herschel Spacecraft key characteristics

S/C Type:	Three-axis stabilised
Operation:	Autonomous (3 hours daily ground contact period)
Dimensions:	7.5 m high x 4.0 m diameter
Telescope diameter:	3.5 m
Total mass:	3170 kg
Solar array power:	1500 W
Average data rate to instruments:	130 kbps
Absolute pointing Error (APE):	1.90 arcsec (pointing) / 2.30 arcsec (scanning)
Relative Pointing Error (RPE, pointing stability):	0.19 arcsec (pointing)
Spatial Relative Pointing Error (SRPE):	< 1.5 arcsec
Cryogenic lifetime from launch:	min. 3.5 years

2.1.1. Herschel Extended Payload Module

The EPLM is mounted on top of the satellite bus, the service module (SVM) and consists of the cryostat containing the instruments' focal plane units (FPU) and the Herschel telescope. The following sections describe the main components of the payload.

2.1.1.1. The Telescope

So that the favourable conditions offered by being in space can be exploited to the full, Herschel carries a precision, stable, low background telescope (Figure 2.2). The Herschel telescope is passively cooled, allowing the size limitations imposed by active cooling to be overcome. Thus its diametre is only limited by the size of the fairing on the Ariane 5-ECA rocket. The Herschel telescope has a total wavefront error (WFE) of less than 6 μm (corresponding to "diffraction-limited" operation at < 90 μm) during operations. It also has a low emissivity to minimise the background signal, and the whole optical chain is optimised for a high degree of straylight rejection. In space the telescope cools radiatively, protected by a fixed sunshade, to an operational temperature in the vicinity of 85 K, with a uniform and very slowly changing temperature distribution.

The chosen optical design is a classical Cassegrain with a 3.5-m diameter primary and an "undersized" secondary. The telescope has been constructed almost entirely of silicon carbide (SiC). The primary mirror (M1) has been made out of 12 segments that have been brazed together to form a monolithic mirror, which was machined and polished to the required thickness (~3-mm) and accuracy. The secondary mirror (M2), with 308-mm diameter, has been manufactured in a single SiC piece. It is adjusted on the SiC barrel by tilt and focus adjustment shims. In order to avoid the Narcissus effect on the detectors, the central part of the secondary mirror is shaped in such a way that no parasitic reflected beam can enter the focal plane.

The hexapod structure (also made of SiC) supports M2 in a stable position with respect to M1. Finally, three quasi-isostatic bipods, made of titanium, support the primary mirror and interface with the cryostat. The focus is approximately one metre below the vertex of M1, inside the cryostat.

The proper telescope alignment and optical performance have been measured on ground in cold conditions. The measured wavefront performance in cold is in line with the requirements. In-flight results confirm the correctness of the focus position.

The M1 and M2 optical surfaces have been coated with a reflective aluminium layer, covered by a thin protective "plasil" (silicon oxide) coating. The telescope was initially kept warm after launch into space to prevent it acting as a cold trap while the rest of the spacecraft was cooling down.

Figure 2.2. The Herschel telescope flight model.

Key telescope data are summarised in Table 2.2.

Table 2.2. The Herschel Telescope's predicted characteristics at working temperature (70 K)

Configuration:	Cassegrain telescope
M1 Free diameter:	3500-mm
Focal length:	28500-mm
f-number:	8.68
Field of View radius:	0.25°
M1 curvature radius / conic constant:	3499.02-mm / -1
Aperture stop / distance to M1 apex:	M2 mirror / 1587.555-mm
M2 diameter:	308.11-mm
M2 curvature radius / conic constant:	345.2-mm / -1.279
Image diameter:	246-mm
Image curvature radius / conic constant:	-165-mm / -1
On-axis best focus distance to M1 vertex:	1050-mm

2.1.1.2. The Cryostat

The Herschel cryostat houses the focal plane units of the three scientific instruments depicted in Figure 2.3. The cooling concept for the Herschel instruments is based on the proven principle used for the ISO mission. The temperature required in the instrument focal plane is provided down to 1.7K by a large superfluid helium Dewar (helium at 1.6K), sized for a scientific mission of 3.5 years. This is achieved with a total amount of 2160 litres of helium cryogen. The cryostat provides 1.7K as its lowest service temperature to the instruments. Further cooling down to 0.3K, required for two instruments (the SPIRE and PACS bolometers) is achieved by dedicated ^{3}He sorption coolers that are part of the respective instrument focal plane unit. In orbit the liquid Helium is maintained inside the main tank by means of a phase separator (a sintered steel plug). The heat load on the tank will evaporate the Helium over the mission time at an estimated rate of about 200 grams per day. The enthalpy of the gas is used efficiently to cool parts of the instruments that do not require the low temperature of the tank (two temperature levels, at around 4K and around 10K). After leaving the instruments the evaporated gas is further used to cool the 3 thermal shields of the cryostat.
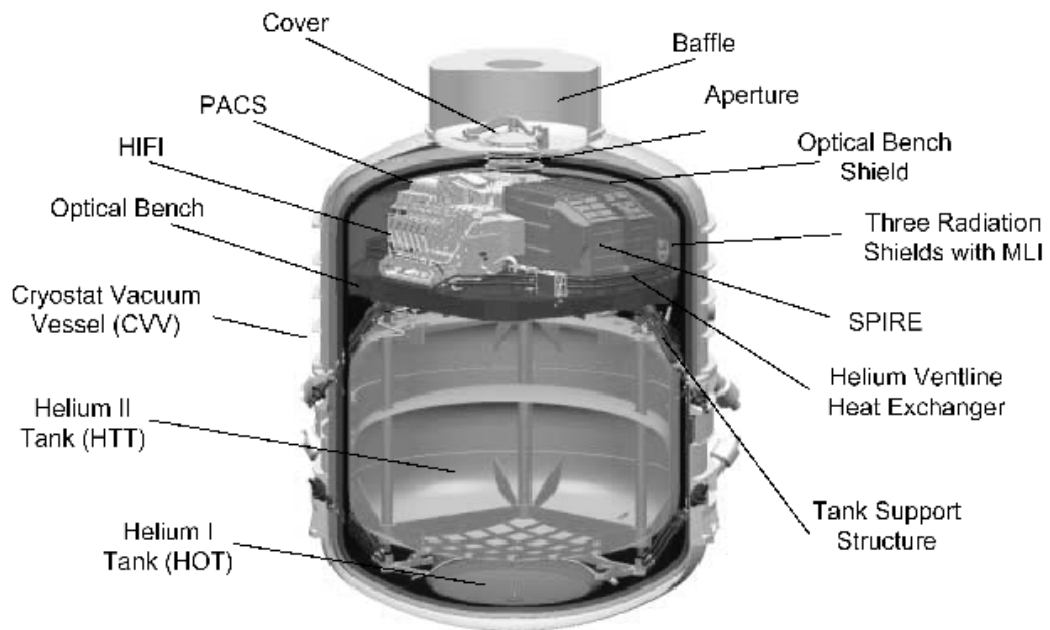
Figure 2.3. The Herschel cryostat.

During ground operations, the vacuum vessel was closed by the means of a cover, located at its top, which was opened once in orbit. To maintain a cold environment inside the cryostat during the last few days before launch in Kourou, an auxiliary liquid Helium tank was used. The space side of the Cryostat Vacuum Vessel (CVV) is used as a radiator area to cool the CVV on orbit to a final equilibrium temperature of about 70K. This radiator area is coated with high emissive coating to achieve low temperatures in the L2 orbit. Multi-Layer-Insulation (MLI) covers the outer CVV-surfaces, in order to insulate it from the warm items (satellite bus and Sunshield). The outer layer of the MLI is optimised for the lowest temperature of the CVV. The outside of the cryostat is the mechanical and thermal mounting base for the Herschel telescope, the local oscillator unit of HIFI, the Bolometer Amplifier Unit of PACS and the large sunshield protecting the CVV from the sun.

2.1.1.3. Instruments

The science payload is accommodated both in the "cold" (CVV) and "warm" (SVM) parts of the satellite. The instrument FPUs are located in the "cold" part, inside the CVV mounted on the optical bench, which is sitting on top of the superfluid helium tank. They are provided with a range of interface temperatures from about 1.7 K by a direct connection to the liquid superfluid helium, and additionally to approximately 4 K and 10 K by connections to the helium gas produced by the boil-off of liquid helium gas, which is used efficiently to provide the thermal environment necessary for their proper functioning. The "warm" - mainly electronics - parts of the instruments are located in the SVM. The following instruments are provided within the Herschel spacecraft:

- The Photodetector Array Camera and Spectrometer (PACS)

- The Spectral and Photometric Imaging REceiver (SPIRE)

- The Heterodyne Instrument for the Far Infrared (HIFI)

The instruments are described in their respective users' manuals

2.1.2. The Service Module (SVM)

The service module (SVM) is the box-type enclosure at the bottom of the satellite, below the EPLM and carries all spacecraft electronics and those instrument units that operate in an ambient temperature environment. It is depicted in Figure 2.4.
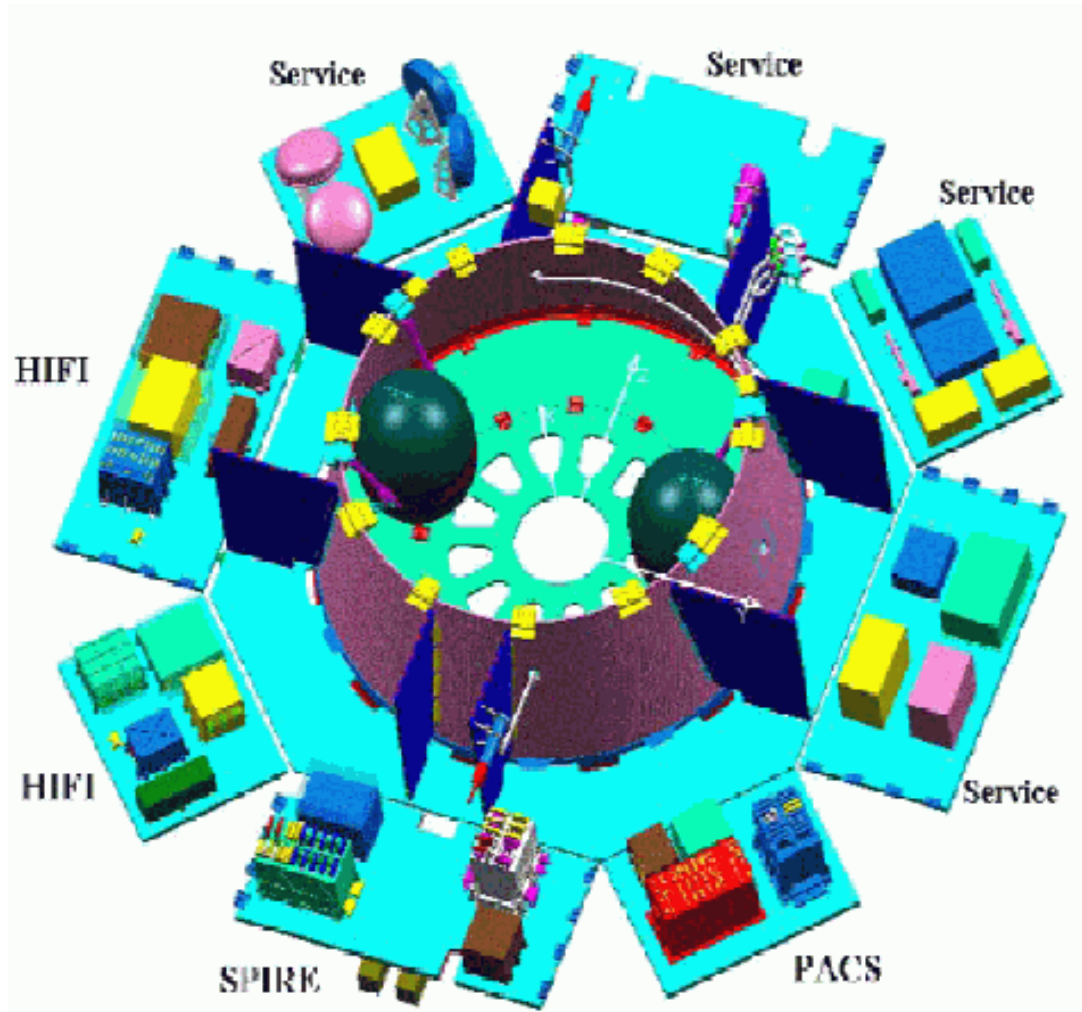
Figure 2.4. The Herschel service module.

SVM modularity is achieved by implementing units of similar function on each of the panels. Panels are either dedicated to one instrument or to a single sub-system (Attitude Control, Power, Data handling-telecommunications). The propellant tanks are symmetrically implemented inside the central cone. The SVM also ensures the mechanical link between the launcher adapter and the EPLM.

2.1.2.1. The Sun shield and solar arrays.

The electrical power of the satellite is produced by the solar array. The solar array is in front of the cryostat to protect it from solar radiation. The rear of the sunshield is covered with multi layer insulation as is the part of the cryostat facing this warm part of the system. The geometrical design has to consider the size of the cryostat and the telescope, the required sun aspect angles of the s/c in orbit and the limited diameter of the fairing of the launcher. For Herschel a relatively simple system with a fixed solar array has been selected. The lower part actually carries the solar cells. The upper part is free of solar cells to allow it to be at a lower temperature, which in turn helps for the telescope to stay at the required temperature. The height of the sunshield is driven by the need to shade the entire telescope when the spacecraft is pointed closest to the sun (60° Sun aspect angle).

2.1.3. Spacecraft Axes definition.

The Herschel s/c coordinate axis system is defined in [RD1] as follows:

- The positive X-axis is perpendicular to the separation plane and nominally coincides with the

longitudinal launcher axis. The positive X-axis shall be along the nominal optical axis of the Herschel telescope, towards the target source.

• The Z-axis forms a plane with the X-axis perpendicular to the separation plane such that nominally the Sun lies in the XZ plane (zero roll angle), positive towards the Sun. In other words, the XZ plane is the symmetry plane of the solar array, the Z-axis pointing outwards from the solar array.

• The Y-axis completes the right-handed orthogonal reference frame.
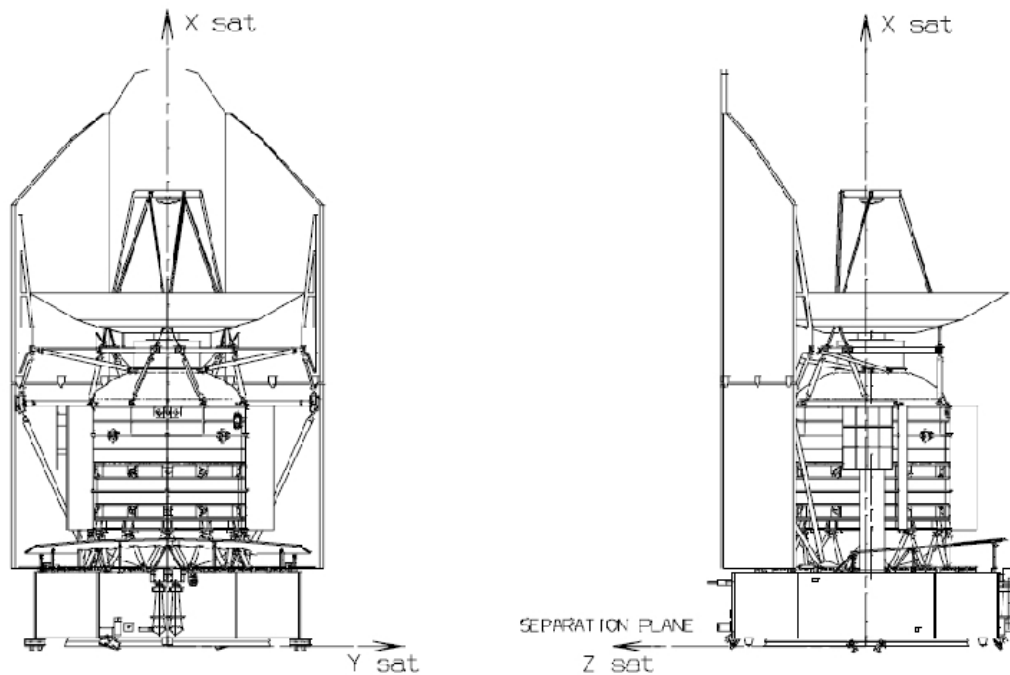
Figure 2.5. Herschel s/c axes (from [RD1])

2.2. Spacecraft orbit and operation

Herschel and Planck was launched aboard a single Ariane V ECA launch vehicle from European spaceport at Kourou. The launch made use of the Sylda 5 adapter with Planck being the lower passenger below the Sylda 5 and Herschel mounted as upper passenger. The two spacecraft separated within 30 minutes after launch and proceeded independently to different orbits about the second Lagrange point of the Sun-Earth system (L2). (see Figure 2.6). Even though both satellites orbit around L2, their orbits are quite different. Herschel acquired its final orbital position at around 1.5 million km from the Earth with only a minor correction manoeuvre after a transfer of about sixty days.
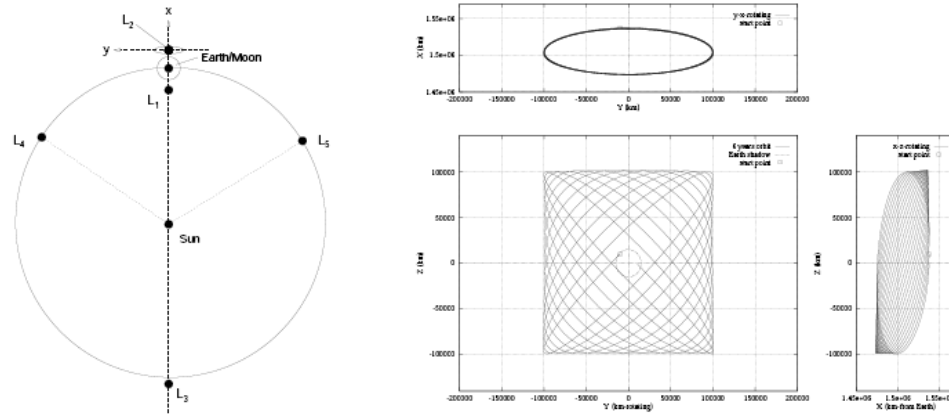
Figure 2.6. *Left:* **Position of the Lagrange points for the Sun-Earth/Moon system. L2 lies 1.5 million kilometres from Earth.** *Right:* **An example of a Lissajous orbit around L2. The orbit x and y-axis are as shown in the plot on the left, the z-axis is normal to paper.**

The Herschel spacecraft was eventually placed in a large "halo" orbit around L2 (halo orbits are special cases of Lissajous orbits around Lagrange points where the in-plane and out-of-plane frequencies are the same), with an amplitude of about 700 000-km and a period of approximately 178 days. The distance from the Earth ranges from 1.2 to 1.8 million km.

The orbit chosen for Herschel presents a number of advantages summarised below:

- Simplifies long observations, since the Sun and the Earth remain close to each other as seen by the S/C (Sun-S/C-Earth angle always < 40°)

- Very stable thermal and radiation environment

- No trace of atmosphere

- A large halo orbit can be achieved without any injection Δv

Major drawbacks are the long distance for communications and the fact that orbits around the L2 are unstable; without orbit corrections the spacecraft would deviate exponentially from the nominal one. Small correction manoeuvres, applied at approximately monthly intervals, maintain the orbit close to the nominal one. Figure 2.7 shows an example of large halo orbit around L2 (from [RD2]).
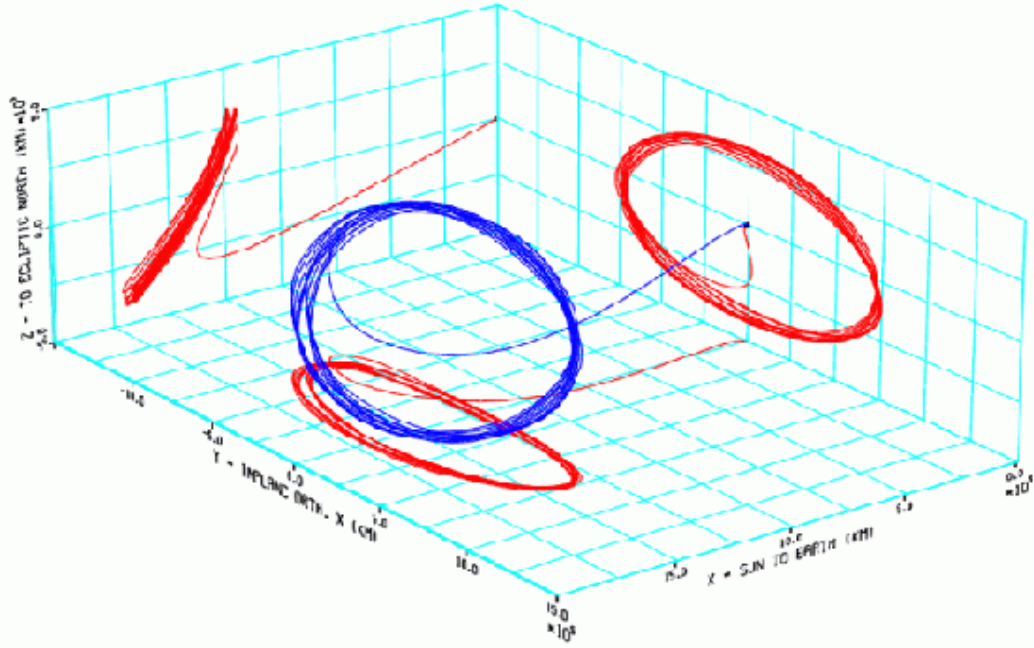
Figure 2.7. A 3D representation of a large halo orbit around L2. The Earth is located at (0,0,0). Red tracks are the projection on the three orthogonal planes of the 3D orbit (blue track).

Herschel operations are performed by the European Space Operations Centre (ESOC) located in Darmstadt (Germany). The main ground station is New Norcia (Australia), which is equipped with a a 35-metre antenna using X band up and down links. New Norcia is backed up by the Cebreros ground station (Spain). In the phase immediately after launch the Kourou (French Guiana) and Villafranca (Spain) ground stations were also used. During routine operations, the ground station communication link is restricted to a duration of approximately 3 hours. During this time, the spacecraft antenna is be pointed to the Earth. The data stored in the on-board solid state mass memory are downlinked, and the mission time line with the new schedule is uplinked. Real time operations and spacecraft maintenance are also carried out during this period. The rest of the time the satellite operates autonomously. The system has been designed to support 48 hours of autonomous operation, with requires a solid state mass memory capability of 25 Gbt. The amount of Herschel data downloaded per day is in excess of 8 Gbt.

2.3. Sky visibility

The areas of the sky accessible to the Herschel telescope are determined by a number of constraints applicable to Sun, Earth, Moon and other bright solar system objects. In particular, the following constraints are applicable through the mission:

- Sun-S/C-LoS angle in the S/C XZ plane (Solar Aspect Angle or SAA) of 60°.8 to 110° for normal operations. Please notice that the allowed range has been reduced with respect to the original one (60° to 120°) since in the extreme SAA range ('warm' attitude range, SAA in the 110° to 120° interval) a noticeable pointing performance degradation (larger APE and pointing offset drift) due to thermo-elastic effects has been observed. Moreover, this degradation persists even if the S/C is brought back to 'cold' attitude until the structure settles back in the original position. Nevertheless, if deemed unavoidable, short (less than 1 hour) observations in the 'warm' SAA range (110° to 120°) can be scheduled at the end of the operational days. The observer should be aware that in such cases, a degradation of the pointing accuracy is very likely. Similarly, post launch a small change was made to the extreme range of permitted of solar aspect angles to limit it to a maximum range from 60°.8 to 119°.2.

- Maximum roll angle of ±1°

In addition, the following extreme Earth and Moon angles do occur across the mission (to be taken into account for straylight considerations):

- Sun-S/C-Earth angle of 37°

- Sun-S/C-Moon angle of 47°

In order to avoid straylight pollution and also for safety reasons (to prevent large fluxes of light from reaching detectors), the nominal half-cone exclusion angles listed in Table 2.3 apply to observations towards major planets.

Table 2.3. Nominal exclusion angles (half-cones) for observation towards major planets

Instrument	Mode	Mars	Jupiter	Saturn	Instrument Critical
SPIRE	Slew	15 arcmin	15 arcmin	15 arcmin	Yes[a]
	Pointing	1.5 deg	1.5 deg	1.5 deg	Yes[a]
HIFI	Slew	36 arcmin	36 arcmin	36 arcmin	No[c]
	Pointing	36 arcmin	36 arcmin	36 arcmin	No
PACS	Slew[d]	4 arcmin	4 arcmin	4 arcmin	No
	Pointing[e]	1.5 deg	1.5 deg	1.5 deg	No
a. SPIRE has determined that, while Jupiter and possibly Saturn will not damage the instrument, they would render it inoperable for a significant period (possibly even an entire OD)					
b. For SPIRE PACS parallel mode both the SPIRE and PACS restrictions apply.					
c. HIFI wishes to avoid straylight pollution when observing fainter objects with a SSO close to the instrument LoS. The instrument will not be harmed by the presence of a major SSO in the FoV and will, in fact, even use Mars as its primary calibrator.					
d. During slews, the detectors are ON (photometry, spectroscopy or parallel mode).					
e. During non-SSO PACS observations. PACS may well wish to observe these SSOs directly.					

The time windows when a fixed or moving target or list of targets are visible can be calculated with HSpot. The tool provides an easy way to check in which time intervals a source is visible during the mission. The visibility calculation does not yet take into account the avoidance cones around Jupiter, Saturn and Mars described above.

The sky visibility for each date has been determined by the launch date (14th May, 2009) and the orbit of the satellite. Considering a nominal duration of the operations, all areas in the sky are visible at least 30% of the time. The sky visible region moves slowly on a daily basis. The two snapshots at the bottom of Figure 2.8 illustrate the typical sky visibility differences after a 3 month interval; although this is calculated for an different launch date to the actual one, the graphic remains a valid representation of the effect.
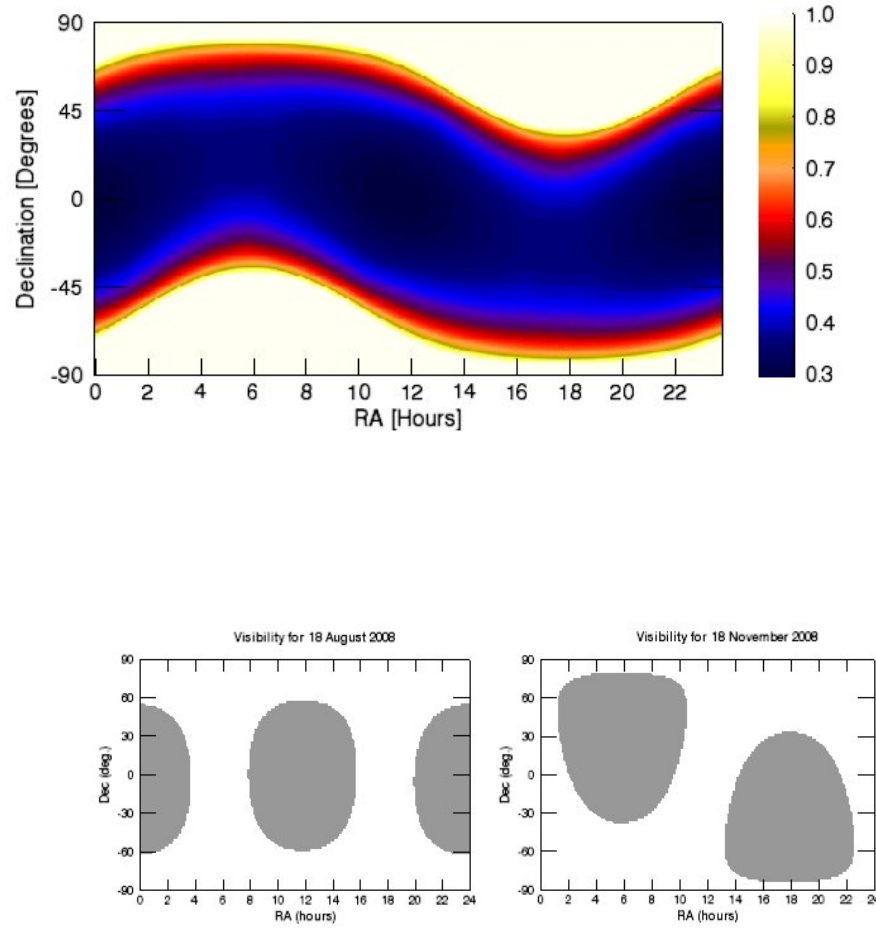
Figure 2.8. Top: The sky visibility across the sky as a fraction of the total hours through the Herschel mission, represented as a colour scale (shown at right) where black represents 30% visibility and white represents permanent sky visibility. Bottom: sky visibility for two sample dates. Shadowed areas represent inaccessible sky areas.

2.4. Herschel pointing performance

This section deals with the pointing performance of the Herschel spacecraft. The spacecraft Attitude Control and Measurement System (ACMS) consists of several components, as depicted in Figure 2.9. The main constituents of the ACMS are the attitude control computer (ACC), gyroscopes (GYR), star trackers (STR), reaction control system (RCS), reaction wheel assembly (RWA), Sun acquisition sensors (SAS), coarse rate sensors (CRS) and attitude anomaly detectors (AAS).
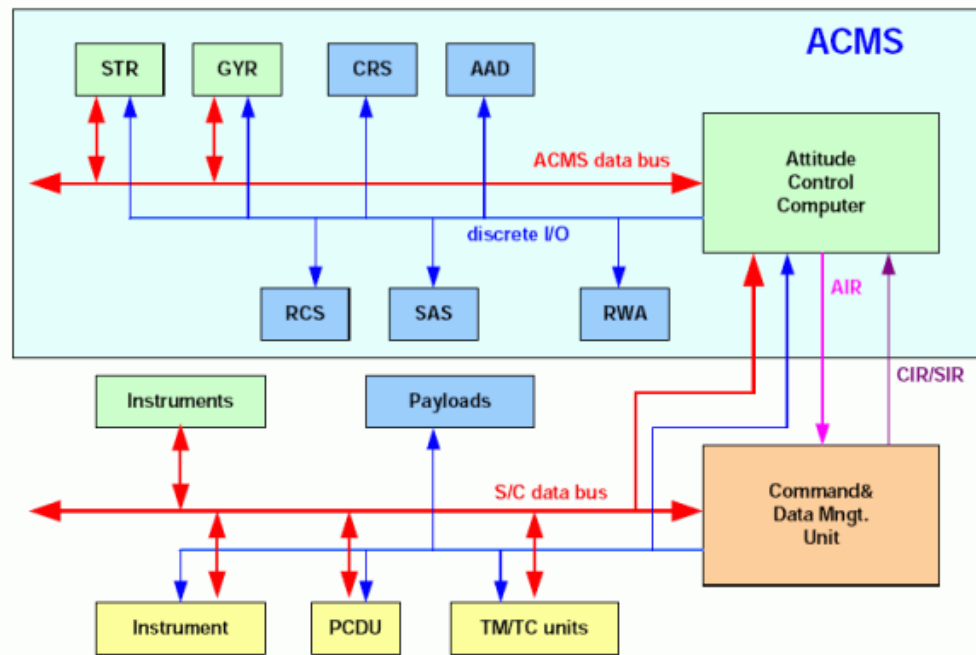
Figure 2.9. Diagram of the Herschel/Planck avionics.

In normal operation, the spacecraft attitude is commanded by means of the reaction wheel system. It comprises four 8.6 kg wheels in a skewed configuration, each with a momentum storage capacity of 30 Nms and a maximum delivered reaction torque of 0.215 Nm in either positive or negative direction.In the baseline configuration, all four whels are powered and used for actuation, providing optimum slew performance and momentum storage. Nevertheless, the ACMS is also capable of operating with only three reaction wheels powered. In the nominal configuration, the maximum slew speed is 0.00204 rad/sec, i.e. ~ 7 arcmin/sec.

In normal science operation, the spacecraft attitude is controlled by means of two components: the star trackers (STR) and gyroscopes (GYR). The STR comprises two cold-redundant units, nominally aligned with the -X axis. The STR hardware include:

- An objective lens.

- A baffle to protect from undesired straylight from the Sun and other bright sources.

- The focal plane assembly, containing a CCD detector and a thermo-electric cooler for CCD cooling.

- The sensor electronics.

From a functional point of view, the STR can be seen as a video camera plus an image processing unit that, starting from an image of the sky, extracts the attitude information measured with respect to the J2000 inertial reference system and delivers it to the ACC. A CPU (ERC32 microprocessor) controls the CCD sensor and also carries the image processing task.

Key characteristics of the Herschel's STR are:

- The ability to determine the inertial position from "lost in space".

- FoV: 16.4×16.4 deg².

- An onboard catalogue, based on Hipparcos, of some 3000 bright stars.

• A minimum of 3 stars, 9 is the maximum due to HW limitations.

The STR bias is the largest contributor to absolute pointing error and is pixel-dependent (some 0.8" × √2)

The STR is provided with an enhanced performance mode the so-called "interlaced mode", only applicable if there are ≥ 15 stars in FoV. The STR samples at twice the nominal frequency (4 Hz), 9 stars at a time. In order to get the maximum accuracy it is necessary that the ACC provides as input to the STR an accurate value of the S/C angular rate (the maximum performance is achieved with rate errors below 0.2 arcsec/sec).

Gyroscopes (GYR) are devices that use a rapidly spinning mass to sense and respond to changes in the inertial orientation of it spin axis. *Rate/rate-integrating gyros* provide high-precision measures of the the spacecraft angular rate. The Herschel's ACMS is provided with four gyroscopes mounted in a tetrahedral configuration. The four gyroscopes are hot-redundant, and each of the four can replace any of the others. The fourth gyroscope is not used for control, but serves to detect an inconsistency in the output of the other three.

The STRs provide an absolute reference, but with limited accuracy. On the other hand, GYRs are very accurate, but only on short temporal (bias drift, 0.0016 deg/hour) and spatial (variation in the scale factor should be taken into account for distances larger than 4 deg) scales. Therefore, the GYR attitude must be recalibrated using the STR information. Therefore, in normal operation the spacecraft attitude is computed by combining the STR and GYR measurements in the ACC using a linear Kalman filter. The so-called "filtered attitude" is sampled and downloaded with a frequency of 4Hz.

Herschel pointing modes are based either on stare pointings (fine pointing mode) or moving pointings at constant rate (line scan mode). Raster maps are 'grids' of stare pointings at regular spacings; in the position switching and nodding modes, the boresight switches repeatedly between two positions in the sky. Scan maps are sequences of line scans at regular spacing. Allowed angular speed ranges from 0.1 arcsec/sec to 1 arcmin/sec. In addition, the Herschel spacecraft can track moving Solar System targets at rates up to 10 arcsec/min.

2.4.1. Pointing accuracy definitions

In this section, formal definitions of the spacecraft pointing accuracy parameters are provided. The term 'pointing', when applied to a single axis (e.g. the telescope boresight), refers to the unambiguous definition of the orientation of this axis in a given reference frame. When characterising the pointing performance of the telescope, it is possible to provide a figure of the absolute attitude accuracy provided by the ACMS (absolute pointing error), or how accurate the 'a posteriori' knowledge of the absolute attitude (the absolute measurement error) can be, or how stable the pointing is (the relative pointing error). Furthermore, the pointing performance can be also characterised in terms of the relative accuracy of a set of attitude measurements (the spatial relative pointing error). The latter measurement is important to characterise the accuracy of the relative astrometry in a map comprising several pointings (e.g. from a raster pointing).

Herschel pointing accuracy definitions, presented below, are based on the prescriptions given in the *ESA Pointing Error Handbook* (ESA-NCR-502):

• **Absolute Pointing Error (APE)**: the angular separation between the desired direction and the actual instantaneous direction.

• **Absolute Measurement Error (AME)**: the angular separation between the actual and the estimated pointing direction (*a posteriori* knowledge).

• **Pointing Drift Error (PDE)**: the angular separation between the average pointing direction over some interval and a similar average at a later time.

• **Relative Pointing error (RPE) or pointing stability**: the angular separation between the instantaneous pointing direction and the short-time average pointing direction at a given time period (in this case 60 sec).

- **Spatial Relative Pointing Error (SRPE)**: angular separation between the average orientation of the satellite fixed axis and a pointing reference axis, which is defined to an initial reference direction.

2.4.2. Pointing performance

The main pointing error contributors within the Herschel spacecraft are:

- To AME and APE:

  - Position-dependent bias within STR. *It is also the main contributor to SRPE.*

  - Residuals from calibration

  - Thermo-elastic stability of the structural path between STR and FPU

  - Instrument LoS calibration accuracy w.r.t. ACA frame (best for PACS)

- To PDE: Thermo-elastic stability

- To RPE: The main contributor is the noise in the control loop comprising STR+Gyro noise attenuated by a linear Kalman filtering.

Table 2.4 summarises the pointing performance of the Herschel spacecraft. The most outstanding non-compliance is related to the SRPE (required 1 arscec vs. predicted/measured performance 2.44/1.45 arcsec).

Table 2.4. Herschel pointing requirements (from SRS v3.2) compared with predictions and measured performance. Goal conditions assume 18 stars available for guidance within the STR

	Baseline (arcsec)		Goals (arcsec)	
Name	Requirement	Performance	Requirement	Performance
		Predic./Measur.		Predic./Measur.
APE point	3.7	2.45/1.90	1.5	1.45/1.35
APE scan	3.7	2.54/2.30	1.5	1.63/n.a.
SRPE	1.00	2.44/1.45[*]	1.00	1.52/1.1

[*]The SRPE has been only measured for small (1 arcmin) two-point rasters.

2.4.3. Gyro propagation mode

As commented above, the STRs provide an absolute reference, but are not accurate enough on their own to satisfy the performance requirements. In particular, they are responsible for the SRPE non-compliance. GYRs only produce accurate attitude measurements in short temporal and spatial scales and their measurements should be recalibrated using the STR information. A mechanism has been devised to perform SRPE-compliant raster pointings by using exclusively the accurate gyro information. Two variants of this mechanism can be considered:

- On-board gyro-propagation mode or Calibration Pointing (CP). This procedure is implemented within the ACMS software *only for the basic raster mode*; gyro-propagation is performed on-board. The gyro-propagated attitude estimates are provided in S/C housekeeping telemetry.

  Warning
  At the time of writing these lines (March 2011), the use of this mode in PACS spectroscopy observations is being assessed.

- On-ground attitude reconstruction by gyro-propagation. This is a ground procedure implemented within the FDS software that reconstructs the attitude estimates based on rate information provided by the gyroscopes. It is intented to improve our *a posteriori* knowledge of the S/C attitude. It is applicable to any mode with OFF positions (i.e. nodding, raster with off position, line scan with off position).

> **Warning**
> The performance of the on-ground attitude reconstruction by gryo-propagation is below the expectations (in general there is no noticeable improvements with respect to the 'standard' filtered attitude etimates) and therefore is not offered as a common-user functionality. For specific enquiries about this topic, please contact Helpdesk http://herschel.esac.esa.int/esupport/.

If gyro-propagation is to be used within an operational day (OD), the following steps must be considered:

- Once per OD, an initial fixed pointing of about 60 min is made to calibrate the GYR bias. Whenever gyro-propagation is requested, this is taken into account and a slot is included within the DTCP.

- Within the next science window period (i.e. the rest of the OD), gyro-propagation observations can be scheduled, provided that they respect the following conditions:

  - An initial 300 sec calibration in the observation gyro calibration position (GCP, a.k.a. OFF position)

  - <600 sec between the recalibrations of the GYR

  - 60 sec periodic recalibration at the GCP.

Chapter 3. Overview of scientific capabilities

Herschel is a versatile observatory with a wide range of capabilities that cover point-source photometry, imaging, large area mapping and spectroscopy at both intermediate and high resolution. Despite the relatively small size of far-IR detectors compared to their visible and near-IR equivalents, it can map large areas of sky efficiently to faint limits. The telescope was designed to give diffraction-limited images - resolution 6 arcseconds - at 90 microns but, in space, it actually performs significantly better than this, with diffraction-limited images being seen as short as 70 microns, with a FWHM of 5.5 arcseconds at this wavelength.

> **Note**
> In mapping mode, at fast scan speeds there is, as is logical, some degradation of the PSF.

3.1. General aspects

The Herschel Space Observatory covers the wavelength range from 55 - 672 microns. This corresponds to the maximum of emission for black bodies in the range from 5-50K approximately. Hence Herschel is be best suited to observing icy outer solar system objects and cool and cold dust in the universe, both in the rest frame and redshifted. A prime objective has been to study the formation of galaxies in the early universe, as cool dust is an excellent tracer of star formation. The Herschel range is also the one at which cool and cold gases emit their strongest lines, meaning that Herschel is also a superb laboratory for examining the chemistry of planetary atmospheres and of the interstellar medium.
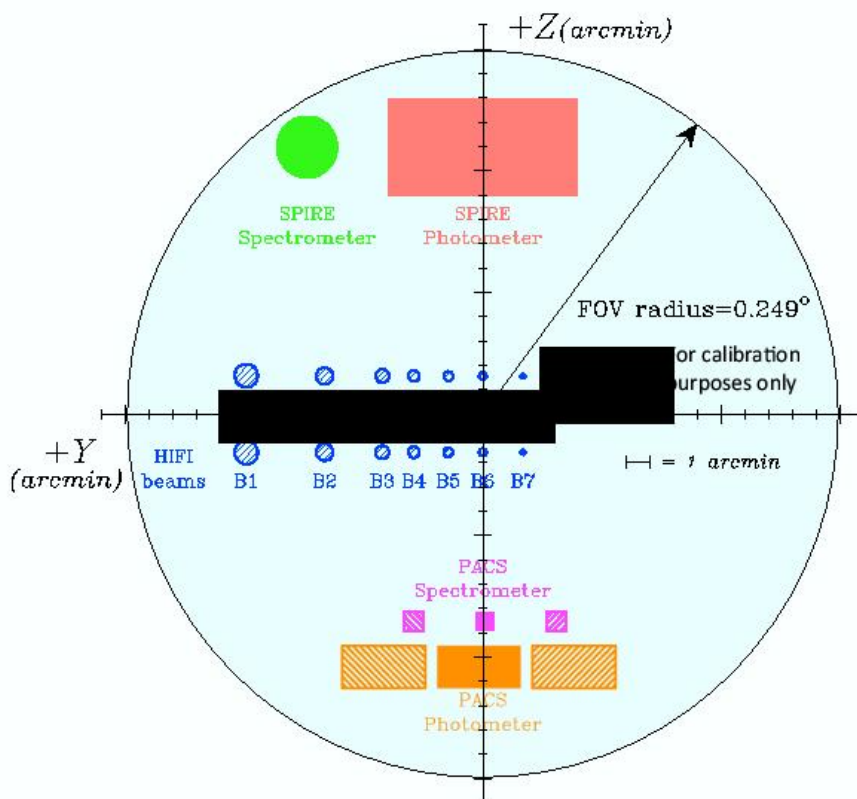

Figure 3.1. The Herschel Focal Plane.

Note
The short wavelength cut-off for Herschel is a matter of definition. For PACS the detector sensitivity below 55 microns is too low to be of practical use and so this value is given as a limit here.

The Herschel Focal Plane is shown in Figure 3.1. The different instrument arrays and apertures are labelled. The full, unvignetted field of view is approximately half a degree.

3.2. Urgent scheduling requests and ToOs

There are a series of factors that decide how quickly Herschel can react to an urgent scheduling request. These are described briefly below. The faster the reaction required, the bigger the effort that is required and the greater the knock-on effects and risk to normal spacecraft operations. In general, any change to the observing schedule made less than 3 weeks before execution of the observations requires special treatment and must be justified carefully. The schedule will only be changed once submitted to MOC if there is a contingency (an instrument problem or operational issue that would lead to a significant loss of observing time, or a ToO); it will not be changed for "routine" tweaking of AORs.

The bottom line is that Herschel can, in normal circumstances, only guarantee to react in 7 days from an urgent request but can react in 5 days in favourable circumstances. Most of the steps described below apply too to normal observations although, of course, on a much less compressed schedule.

The decision to attempt a fast turnaround time for an urgent scheduling request is not taken lightly. MOC at Darmstadt have to confirm that they can process the re-delivered observing schedule in time and the HSC Mission Planning Group have to agree to a delivery schedule that allows MOC enough time to carry out the full processing and check procedure.

3.2.1. Ground station access to Herschel

The Herschel Space Observatory is fundamentally an "off-line" mission and has been designed as such, so reacting rapidly is problematic. "Off-line" means that, instead of being in permanent contact with Herschel there is a Daily TeleCommunications Period (DTCP), normally for 3 hours each day, when the data stored on board must be downlinked and new observations uplinked. The DTCP is within a few hours of local midnight at the ground station (prime is New Norcia, back-up is Cebreros). The DTCP time will also vary according to whether Planck or Herschel is visible to the antenna first - every few months their orbits cross over and the lead satellite in the DTCP changes - and also to demands on Ground Station resources from other missions.

This means that everything must be ready for uplink before the DTCP starts. This DTCP start time defines an unbreachable barrier and everything is calculated backwards from this moment.

At each DTCP, observations are uplinked for execution from 27 to 51h ahead. This means that even if a Ground Station pass is missed due to a communication problem, there are enough observations in the on-board memory for Herschel to continue operating until the end of the DTCP period the following day. All processing time must be added to this minimum 51h.

3.2.2. How is a ToO alert triggered?

There are two methods. Our preferred method is "pre-approval", but ToO alerts may be triggered without pre-approval if the case is compelling, the circumstances justify it and there is a demonstrated need to react quickly to avoid missing a major scientific opportunity.

3.2.2.1. Pre-approval by HOTAC

This is our preferred method. A normal proposal is submitted in the regular Calls for Proposals and evaluated. The proposal should included clear trigger criteria and an expected reaction time that allows its feasibility to be assessed. If HOTAC approves the request, the HSC is required to carry out the observations if the trigger criterion is met and scheduling constraints permit it.

> **Note**
> Sometimes a request may simply not be feasible for operational reasons such as having the wrong instrument active, or because of operational constraints

This method has various advantages, not least of which is that the PI knows in advance that the observations are approved and will be made if possible and will be propietary. The observations are also already available in the HSC database, technically checked and only need to be activated, saving valuable time on activation.

If the PI is satisfied that the trigger conditions are met and that the target is visible, he or she should send a Helpdesk ticket in the ToO Department, supplying all the necessary information to allow the observations to be taken and a short justification of the trigger. The box stating that the ToO is pre-approved by HOTAC should be ticked and the proposal name given to identify it. When the ticket is sent an SMS message will go direct to the Project Scientist and HSC, alerting them of the triggering.

3.2.2.2. Spontaneous ToO alerts

Not all ToOs can be anticipated. Sometimes something will happen that is too good an opportunity to miss; it may be a newly discovered comet that will become bright in 6 months time, or a sudden and unexpected outburst of a known object. In this case things can be more complex and may be slower.

The PI should fill out the same form in the http://herschel.esac.esa.int/esupport/ ToO Department. The alert must be justified and the required observations either provided or, at very least, described in enough detail that an expert at HSC can prepare them. When the ticket is sent an SMS message will go direct to the Project Scientist and HSC, alerting them of the triggering. No action to schedule observations will be taken at HSC until the Project Scientist has approved the request. The Project Scientist's first reaction will usually be to request a technical assessment of the observations and target visibility from HSC. This may require some backwards and forwards iteration with the PI to get the observations right so that the time impact can be assessed and the feasibility ensured. If the request is for observations of a Solar System Object (SSO) it may be necessary for software support at the HSC to add the SSO ephemeris to HSpot. If the requested observing time is significant, the Project Scientist may seek the approval of the HOTAC Chair to add the observations to the schedule. Approval (or not) is then communicated by the Project Scientist in a reply to the Helpdesk ticket.

3.2.3. Processing an urgent scheduling request

Once the observations are approved by the Project Scientist, they are processed. The target day for execution is identified in consultation with MOC. Observations are normally submitted to MOC 3 weeks in advance of execution. This means that any request for an observation less than 3 weeks ahead will require re-processing at HSC and at MOC. Observations already in the database will be linked to the correct instrument control software and time estimator version and released to the Mission Planners who then proceed to de-commit the observing schedule for the day so that it can be modified. Observations not in the database must be submitted as a new proposal first to get them into the database. Once a de-commit of a schedule occurs it is removed completely from the system, meaning that even if the same schedule is subsequently re-instated, it must be completely re-processed at MOC, so a schedule is only de-committed when it is certain that a re-plan is necessary and feasible.

The new observation(s) is/are fitted into the schedule and a revised schedule generated. The revised schedule goes through a list of checks and is then submitted to the instrument team(s) and the Project Scientist for final approval; either may request revision to the schedule as drafted if they identify any problems with it. Only when all parties are satisfied will the schedule be transmitted to MOC. This modifies the sequence of pointing commands for the spacecraft and so has a knock-on effect for the following day too because it will affect the initial state of the spacecraft gyros for the next observing day, meaning that MOC must generate a new solution for obtaining the desired the pointing for both days, even though HSC has only re-generated the schedule for the first of them.

The time required at HSC to process the observations and generate the new schedule depends on the complexity of the changes to the schedule and the availability of the specialists to approve the draft schedule and any modifications requested: it may require a full working day.

When the new schedule is approved, it is transmitted to MOC by FTP. Initial processing is automatic. Once this is complete a member of the Flight Dynamics Team will continue processing interactively. If all the checks are passed, the schedule will be passed on to a member of the MOC Mission Planning Team who will generate the file containing all the telecommands that will be transmitted to the spacecraft. This is usually the slowest and most complex part of the processing chain due to the serious potential consequences of any error. This processing at MOC requires a full working day and must be completed in advance of the DTCP; if the DTCP is early in the day or at the weekend this means that the task must be completed the previous working day.

Essentially, the peculiarities of the Herschel mission compared to on-line missions and the series of effectively uncompressable response times at different stages in the processing chain provide a fundamental limit to reaction times in any re-planning scenario.

3.3. Photometry with Herschel

3.3.1. Instrument capabilities

The full wavelength range of Herschel is covered by six broadband ($\Delta\lambda/\lambda=3$) filters. In SPIRE, all three filters (250, 350 and 500 μm) are imaged simultaneously on three spiderweb bolometer arrays. PACS users are able to image with a "red" (130-210 μm) and a "blue" (either 60-85 or 85-130 μm) filter simultaneously on two bolometer arrays.

Note
Colloquially, but inaccurately, the 130-210, 85-130 and 60-85 micron filters are often referred to as the PACS "red", "green" and "blue" bands.

This makes Herschel a superb instrument for multicolour surveys. SPIRE can image a square degree of sky to an instrument noise equal to the extragalactic confusion limit (1 sigma) in 3.3 hours and, of course, much larger areas of the sky to a lesser sensitivity; note that because of the overheads involved in scan maps (i) making large maps is more efficient than making small maps and (ii) making square maps, or maps with a small aspect ratio is somewhat more efficient than making long, thin strip maps.

The main imaging capabilities are summarised in Table 3.1. As dust is a strong tracer of star formation, one of Herschel's greatest strengths is the possibility to study the history of star formation in the universe. By combining PACS and SPIRE data, users will be able to follow the dust emission signature of starbursts redshifted to increasing wavelengths in ever more distant galaxies; this makes Herschel an enormously powerful facility for studying the formation and evolution of galaxies.

Observations with Herschel will give a new insight into the process of star and planet formation. Herschel can study both the processes of star formation in molecular clouds and the debris disks that are the tracer of planetary system formation in young stars. To date, few debris disks are known and observations with Herschel, with its wide wavelength coverage, will allow many more to be detected and studied. Similarly, Herschel observations will be valuable in the study of the later phases of stellar evolution, particularly circumstellar shells, mass-loss in general and stellar winds.

Finally, Herschel is a powerful tool for studying the physics of the more distant and colder objects of the solar system: such as the atmospheres of the giant planets, their icy satellites, cometary nuclei and cometary atmospheres. Herschel observations are permitting the albedos and thus the surface conditions and diameters of these bodies to be measured with great precision.

Table 3.1. The main imaging capabilities of PACS and SPIRE. Please note that the wavelength range of detector sensitivity is approximate and the instrument sensitivities depend on the observing mode, so the values given are only orientative: please consult the relevant observing manual for more detailed values.

	PACS	SPIRE
Wavelength range	60-210 μm	200-670 μm
Field of view	1.75x3.5'	4x8' (unfilled)
Pixel size	3".2 (60-130 μm), 6".4 (130-210 μm)	18".2 (250 μm), 24".9 (350 μm), 36".3 (500 μm)

Typical sensitivity (5σ/1hr)	5 mJy (70/110 μm bands), 10 mJy (160 μm band)	Will reach the confusion limit (see below)
Confusion limit (ideal case)	<0.1 mJy (70 μm), 0.27 mJy (100 μm), 0.92 mJy (160 μm)	5.8mJy (250 μm), 6.3mJy (350 μm), 6.8mJy (500 μm)
Filters	60-85 or 85-130 μm and 130-210 μm (simultaneous)	250, 350 and 500 μm (simultaneous)

3.3.2. Using SPIRE and PACS in parallel

Herschel offers a parallel mode for users who wish to carry out large-scale mapping programmes with a wide range of wavelength coverage.

3.3.2.1. The benefits of using parallel mode

Parallel mode allows observers to use both SPIRE and PACS simultaneously in a fast (60 arcsec/s) and nominal speed (20 arcsec/s) scanning mode to cover very large areas of sky quickly in all three SPIRE bands and in two of the three PACS bands, to a modest sensitivity. This mode is intended to make ambitious, multi-band, large area mapping programmes more efficient than carrying them out individually with each instrument in turn. In this mode SPIRE is the prime instrument and thus the driver in defining observations and PACS data should be treated more as a "bonus" to observers.

3.3.2.2. The limitations of using parallel mode

SPIRE and PACS point at different places on the sky separated by 21 arcminutes. This means that this mode is extremely inefficient at mapping small areas of sky. Although a minimum area of 30x30 arcminutes for a Parallel Mode map is permitted by HSpot, alternatives should certainly be considered for any area of sky smaller than one square degree and possibly even for larger areas than this.

While the SPIRE integration gets to a depth which is fairly close to the confusion limit, the depth of exposure is relatively less for PACS, thus Parallel Mode photometry should not be regarded as an adequate substitute for even moderately deep PACS scan maps.

Although very large areas of sky can be mapped quickly at the high scan speed, the speed of scan in sufficiently high that some telescope movement occurs before detector readout is complete, giving rise to a small degree of PSF smearing in the scan direction.

Parallel Mode entails having both PACS and SPIRE cool simultaneously and is thus more demanding in terms of helium usage. As a rule of thumb, for each three parallel cooler re-cycles (6 days of Parallel Mode use) we reduce mission lifetime by one day with respect to using just PACS or SPIRE alone. For efficiency of helium consumption, using Parallel Mode is discouraged unless genuinely required scientifically.

3.4. Spectroscopy with Herschel

Herschel offers two types of spectroscopic capability. PACS and SPIRE offer low to intermediate resolution spectroscopy covering the full Herschel wavelength range. HIFI offers high-resolution spectroscopy over the range from 157-625 μm (480-1910 GHz) using heterodyne techniques, although there is a small gap in coverage from 213-240 microns (1272-1430 GHz), between the HIFI 5b and 6a sub-bands. Users will thus be able to select a wide range of resolutions from $\Delta\lambda/\lambda=20$ to $\Delta\lambda/\lambda=10\ 000\ 000$ according to the brightness of their source and the science that is required. The main spectroscopic capabilities are summarised in Table 3.2.

In its highest resolution mode Herschel will offer a velocity resolution as high as 0.3km/s. The wavelength range covered by Herschel has many thousands of lines of water, atomic transitions and organic molecules. This will allow Herschel to study the chemistry of the interstellar medium, tracing water and organic molecules in molecular clouds. Herschel will also be able to study the chemistry of solar system bodies such as the atmosphere of Mars and the comas of comets in unprecedented detail.

All three instruments have a mapping capability in spectroscopic mode, even though HIFI's is somewhat limited, although by no means negated, by the fact that its detector has only a single pixel. PACS and HIFI can scan the detectors across the sky, accumulating spectroscopic data along the length of the scan. SPIRE cannot do that, instead it uses a beam-steering mirror to make filled maps. All three instruments can make a raster map in spectroscopic mode. This allows a spectroscopic survey to be made either of a region that has been mapped in imaging mode, such as a cluster of galaxies, or across a known extended source such as a molecular cloud.

Table 3.2. The main spectroscopic capabilities of PACS, SPIRE and HIFI. For more details please check the relevant instrument manual.

	PACS	SPIRE	HIFI
Wavelength range	55-210 μm	194-313 and 303-671 μm	157-213 and 240-625 μm (with gap)
Field of view	47x47"	2.0' (unvignetted)	Single pixel (see below)
Pixel size	9"	17", 29" (varies across the bands)	39" (488GHz), 13" (1408GHz)
Sensitivity (5σ/1hr, point source)	2×10^{-18} Wm^{-2} (130 μm, 1st order), 5×10^{-18} (70 μm, 3rd order). Continuum: 100 mJy (1st order), 250 mJy (3rd order)	$1.0\text{-}2.2 \times 10^{-17}$ Wm^{-2} (high resolution), 40-88mJy (low resolution) [5 sigma/ 1 hr]	"A few" mK (Band 1a) to 100mK (Band 7b), 1 sigma/1hr
Resolution	900-2100 (1st order, 102-210 μm), 1800-3000 (2nd order, 72-98 μm), 2600-5400 (3rd order, 55-72 μm)	20-1000	1000-10^7

Note

For the latest information on instrument sensitivities please check the Herschel website at ht-tp://herschel.esac.esa.int/.

Note also that the PACS sensitivity below 57 microns is very low, although HSpot permits the entry of line observations at shorter wavelengths.

Chapter 4. Space Environment

This section will deal with "space environment" aspects of the mission that affect the noise level and therefore the observatory sensitivity. These include:

- Background, including the telescope, instruments and the celestial background

- Radiation environment (high-energy particles)

- Source confusion (CFIRB and cirrus spatial structure, resolved or partially resolved galaxies)

- Straylight due to sources inside or outside the FoV and to instrumental self-emission

4.1. Background radiation

4.1.1. Telescope background

The Herschel telescope is located outside the cryostat and protected by the sunshade from direct radiation from the Sun. The currently measured telescope temperature is in the range 83-90 K and shows an annual variation due to the changing heliocentric distance of the Earth and thus L2 (see below). At this temperature, even given a low emissivity, the source contribution is almost always only a small fraction of the telescope background. For comparison, the telescope background 'flux' is of the order of 1000 Jy, while that of Uranus is ~ 250 Jy and Neptune ~ 100 Jy. Therefore, a precise characterisation of its behaviour is of critical importance.
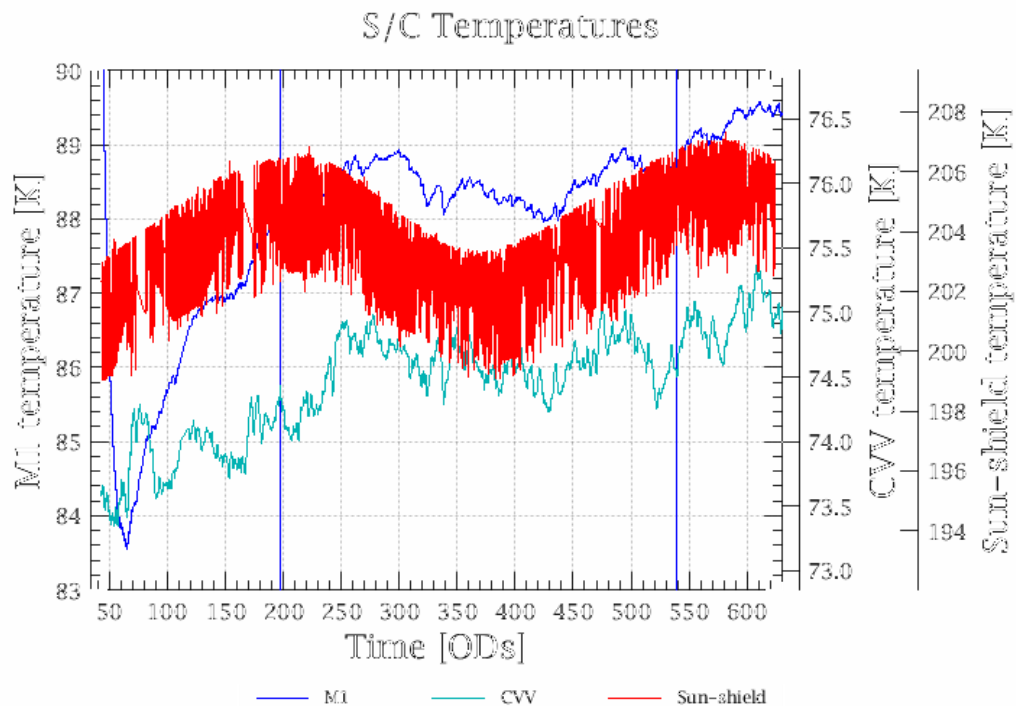
Figure 4.1. Temperatures of the primary mirror (M1), cryostat vacumm vessel (CVV) and sun-shield measured from OD40 to OD660. The monotonic increase of temperature from up to OD300 is well correlated to the seasonal temperature variation model.

The telescope background depends primarily on:

- The average temperature: Figure 4.1 shows the temperature of Herschel's telescope primary mirror measured across 620 ODs (from 22nd July, 2009 to 4th March, 2011; the CVV and sunshield temperatures are also displayed). The data gathered so far indicate that the temperature is some 5K higher than predicted pre-launch (this is compensated by the lower telescope emissivity, see below). On the other hand, the overall trend, characterised by a monotonic temperature increase with a maximum amplitude of some 6K, is well correlated with the 'seasonal' temperature evolution mode.

- The effective emissivity: beyond 100μm, this has a stronger influence on the telescope background level than the temperature. It has been observed that a 1% reduction in emissivity gives a greater improvement than a 5 K reduction in temperature. The predicted (modelled) telescope emissivity was < 0.4%/0.8% (for clean/dusty samples; Fischer et al., 2004 [RD11]). Preliminary results from PACS photometer observations suggest that the actual telescope emissivity is quite significantly lower (about half the predicted figure; Okumura, priv. comm.)

- The straylight (see Section 4.4).

Small spatial and temporal temperature gradients are important to the background stability. The requirements on the primary mirror (M1) are:

- Maximum temperature difference along the S/C Z axis < 10 K (predicted < 0.5 K)

- Maximum temperature difference along the S/C Y axis < 1 K (predicted ~ 0.0 K)

- Along the S/C Z axis: $dT/dt < 13.0$ mK/min

- Along the S/C Y axis: $dT/dt < 1.3$ mK/min

4.1.2. Instruments

See "Self-emission" under Section 4.4

4.1.3. Celestial background

Thermal emission from interstellar dust (known as interstellar cirrus) dominates the FIR Sky Background (FIRSB) at lower Galactic latitudes, while the Cosmic Far-Infrared Background (CFIRB) is more significant towards higher Galactic latitudes, also dominating the confusion noise in the PACS and SPIRE photometric bands. Intrinsically diffuse and unresolved components of the FIRSB are (in descending order of their relative contribution, see [RD4] and [RD5]):

1. Diffuse galactic light (interstellar cirrus): quasi-thermal emission of dust in low density gas clouds in the Milky Way. This is the dominant component for wavelengths λ>70μm.

2. Zodiacal light and emission from the asteroid belt: this is the dominant component of the sky background at MIR wavelengths.

3. Cosmic far-infrared background (CFIRB): accumulated and unresolved light of distant galaxies.

4. The cosmic microwave background (CMB): the CMB also has an important contribution in strength, but the fluctuation amplitudes are small, and well below the detection limits of PACS and SPIRE.

5. Intergalactic diffuse emission

6. Integrated starlight: the integrated contribution from faint stars in the Milky Way is an important component for near- to mid-infrared wavelengths, but has a negligible contribution for longer wavelengths, e.g. those of PACS and SPIRE.

A detailed description of the different components of the FIR background is given in the Herschel Confusion Noise Estimator (HCNE) tool Science Implementation Document ([RD4] and [RD5]). The HCNE tool can be accessed as a standalone service to provide background estimates (see the HSC website for more information) or through the HSpot proposal preparation tool.

While the zodiacal light emission is a major contributor to the sky brightness in the MIR range, it is less important for the FIR and sub-mm wavelengths. Moreover, this emission is quite smooth, lacking fluctuations at arcmin scale (angular resolution of the ISOPHOT instrument on board ISO). Smaller scale fluctuations, in principle, are likely to exist, but the presence of such structures have not been yet confirmed by the recent observations of the Spitzer Space Telescope.

Confusion noise due to the integrated FIR-sub-mm emission from faint asteroids individually below the detection limit has been investigated by Kiss et al. (2006) (see [RD5] and references therein). It has been found that the distribution of asteroids concentrates towards the local anti-solar direction, with a corresponding peak of the confusion noise in the anti-solar point, and an extended cloud is present around the maximum. Seasonal variations are also detected. The confusion noise induced by the cloud of asteroids would only be not negligible in the area around the anti-solar direction, but this area of the sky is closed to Herschel anyway due to the satellite's Sun constraint (see Section 2.3), so the asteroid cloud component is not considered in the HCNE.
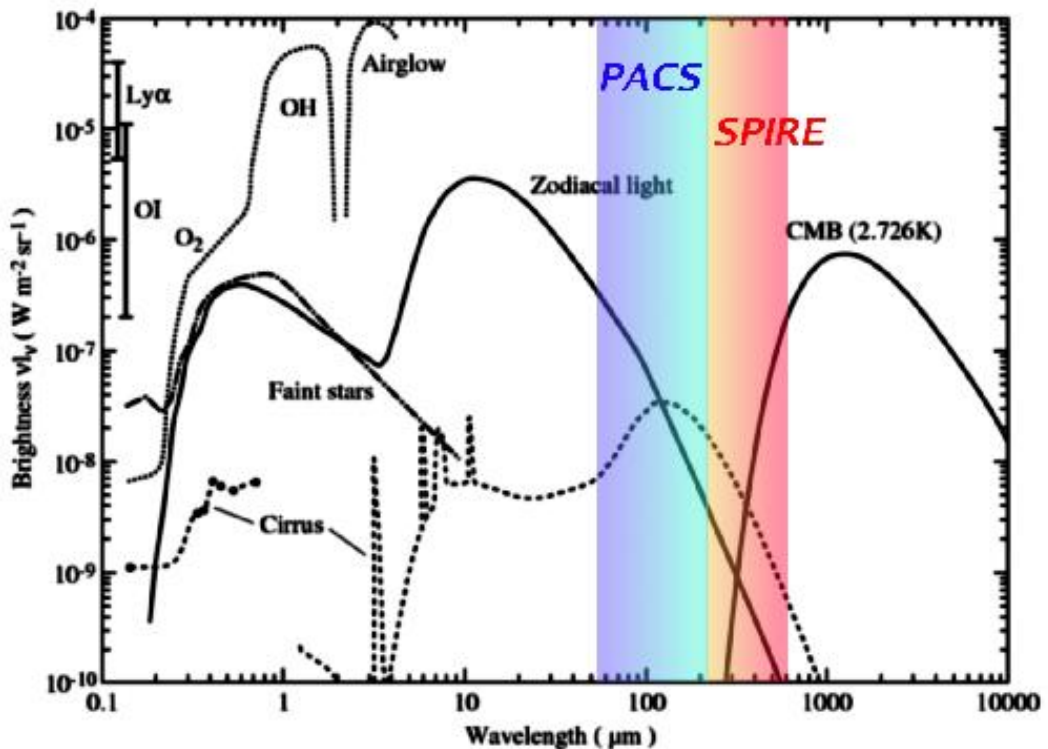
Figure 4.2. Brightness of the night sky, excluding contribution of the extragalactic background (from [RD5], adapted from Leinert et al. 1998, A&A, 127, 1). The spectral range covered by the PACS and SPIRE instruments of the Herschel Space Observatory are indicated. Atmospheric contributors, affecting ground-based observation in the optical and NIR, have been also displayed.

The interstellar medium shows a strong concentration around the Galactic plane; this feature is conspicuous at many wavelengths. However, the cirrus emission is not limited to low Galactic latitudes. It consists of thermal emission of dust in low-density, cool interstellar HI clouds (typically with $T \approx 20K$ and $n \leq 10^2 cm^{-3}$), showing a smooth, modified blackbody SED. It is a strong source of emission, and dominates the sky for wavelengths $\lambda > 70\mu m$, even at high Galactic latitudes. The cirrus emission is highly structured, and shows a typical filamentary structure.

The main characteristic of the cirrus emission is its spatial structure at a specific wavelength. This is usually described by the spectral index, α, of the power spectrum of the image, averaged over annuli

([RD7]). With this parameter the power spectrum is $P=P_o(f/f_o)\alpha$, where P is the power at the spatial frequency f and P_o is the power at the spatial frequency f_o. Due to this parameterisation the structure of cirrus is equivalent to that of a fractal.

According to [RD5], cirrus confusion noise can be generally described by the following equation:

$$\sigma_{cirrus} = c_1 \times (\lambda/D)^{1-\alpha/2} \times B^{\eta}$$

Here σ_{cirrus} is the confusion noise due to the cirrus component, B is the surface brightness of the field α is the spectral index of the logarithmic power spectrum, averaged in annuli (see [RD5] and references therein), λ is the wavelength of the observation and D is the effective diameter of the telescope's primary mirror. The parameters c_1 and η have to be determined from measurements. This is used within the HCNE to compute the noise due to cirrus emission. Details on the computations are given in [RD5].

In many practical cases, Galactic cirrus confusion noise has been found to be easily parameterised as follows (see for instance [RD6] and references therein):

$$\sigma_{cirrus} \sim 0.3(\lambda_{100})^2(D_m)^{-2.5}\langle B_\lambda \rangle^{1.5}$$

where σ_{cirrus} is given in mJy, λ_{100} is the wavelength ratio $\lambda/(100 \ \mu m)$, D_m is the telescope diameter in metres and $\langle B_\lambda \rangle$ is the sky brightness in MJy/sr. If we consider fiducial values $\langle B_{70} \rangle = 0.12$ MJy/sr and $\langle B_{160} \rangle = 1.5$ MJy/sr (corresponding to $N_{HI} = 10^{20}$ cm^{-2}) and $D_m = 3.5$, we get that $\sigma_{cirrus}(70 \ \mu m) = 0.22 \ \mu Jy$ and $\sigma_{cirrus}(160 \ \mu m) = 0.08$ mJy.

4.2. Radiation environment

The L2 environment (and orbits around it) is relatively benign compared to those in geostationary (GO), or low Earth (LEO) orbits. In particular, a series of common threats for satellites in GO or LEO, including the neutral thermosphere, space debris, geomagnetically trapped particles and large temperature gradients, are not a concern for L2 orbits. Environmental aspects to be considered at L2 include:

- Solar wind plasma. Essentially a neutral or cold plasma: 95% protons, 5% He++ and equivalent electrons; 1-10 particles/cm^3. The main risk associated is a low surface charging potential. This plasma may be relatively benign at L2 compared to that found at GO and LEO.

- Ionising radiation: solar flares (energetic electrons, protons and alpha particles), Galactic cosmic rays and Jovian electrons.

- Magnetic fields: Earth's magnetotail extends up to 1000 Earth's radii, so it must be considered (2-10 nT) along with interplanetary magnetic field (~ 5 nT). The effects on the spacecraft and PLM include possible orbit disturbance and electrostatic discharge (ESD).

Therefore, the main radiation components at L2 consist of: Galactic cosmic rays, solar particle events and solar and Jovian electrons. Solar activity follows an 11-year cycle. The last minimum occurred in December 2008 and therefore the Herschel launch in 2009 was carried out during a low activity state. Contrary to former predictions, the next solar cycle will be below average in intensity, with a predicted maximum sunspot number of 90. Given the predicted date of solar minimum and the predicted maximum intensity, solar maximum is now expected to occur in May, 2013 (Solar Cycle 24 Prediction Panel agreement on May 8, 2009). Therefore Solar particle events are expected to be problematic only towards the end of the mission. At the time of writing the AO-2 version of this manual (March 2011), the solar activity is increasing, the proton event observed in OD 663 (7-8 March 2011) being the strongest one observed to date. Notably, there is no noticeable impact of this increased solar activity in the performance of the instruments or the quality of the science products at the levels of activity observed so far (maximum ~50 proton flux units in the March 7/8 event).

Weekly, calibrated plots of the Herschel SREM data and special plots of any observed proton events are available to users, provided by the SREM PI, Petteri Nieminen. These are available at http://proteus.space.noa.gr/~srem/herschel/.

In the early stages of the mission, the dominant radiation source has been Jovian electrons, characterised by a energetic population and a 13-month synodic year modulation. Solar electrons will be an important source at lower energies with abrupt peak emissions, and a 27-day period.

The Herschel spacecraft is equipped with a standard radiation environment monitor (SREM) placed in the -Z SVM panel; the SREM is a particle detector developed for satellite applications that has been added to Herschel and Planck as a passenger. It measures high-energy electrons (from 0.5 MeV to infinity) and protons (from 20 MeV to infinity) of the space environment with an angular resolution of some 20 degrees, providing particle species and spectral information. The SREM data are received on-ground and processed by the Space Weather Group at ESTEC, providing valuable information on the radiation environment at L2. A sample plot showing the calibrated count rates in three counters (TC1 - protons with E > 20 MeV; TC2 - protons with E > 39 MeV; TC3 - electrons with E > 0.5 MeV) is displayed in Figure 4.3 .
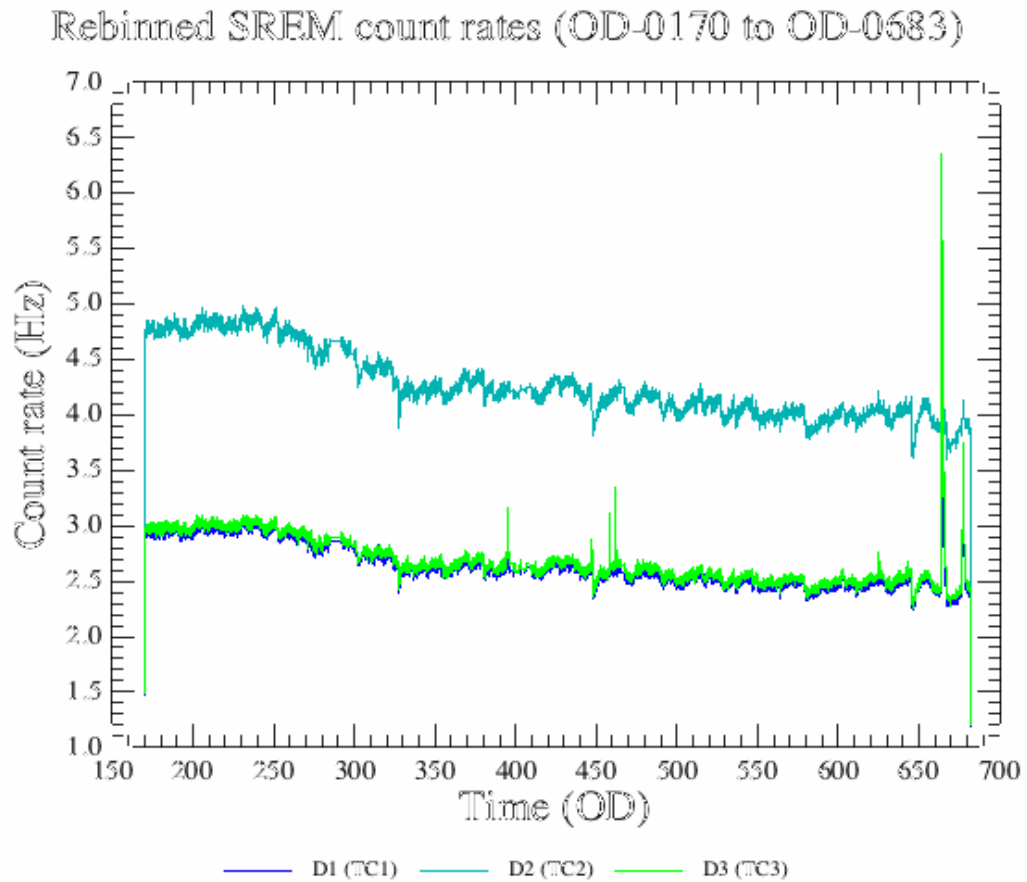
Figure 4.3. SREM calibrated count rates in three counters (TC1, TC2 and TC3), rebinned in intervals of five minutes. from the 30th of October 2009 (OD 170) to 27th March of 2011 (OD 683). The slight decline of the count rates can be explained by an increased solar activity and the subsequent increase of shielding to Galactic cosmic rays. Several events are visble, the most conspicuous a proton flare detected in OD 663 (7-8 March 2011).

4.3. Source confusion

Source confusion is an additional noise factor closely related to the astronomical background, described in Section 4.1. The sensitivity limit due to confusion is determined by the telescope aperture, observation wavelength and the position on the sky. The sensitivity cannot be improved by increasing the integration time after reaching the confusion limit. The most important contributions to source confusion are:

- Structure of the CFIRB, as well as resolved and partially resolved extragalactic sources dominate at high galactic latitudes.

- Small-scale structure in cirrus clouds may dominate at intermediate Galactic latitudes. The contribution depends heavily on the level of cirrus emission at the position on the sky.

The confusion noise is usually defined as the (stochastic) fluctuations of the background sky brightness below which sources cannot be detected individually. In addition to the diffuse Galactic foreground cirrus component, these fluctuations are caused by intrinsically discrete extragalactic sources in the beam. Due to the limited telescope diameter compared to the wavelength, these fluctuations play an important, if not dominant, role in the total noise budget in extragalactic surveys carried out in the MIR, FIR and sub-mm range. Moreover, the noise due to extragalactic sources depends strongly on the shape of the source counts at a given wavelength.

There are two different criteria to derive the confusion noise, and thus the detectability of a point-like or compact source:

- First, the target source flux should be well above the average background fluctuation amplitude. This is the basis of the "photometric criterion", derived from the fluctuations of the signal due to sources below the detection threshold S_{lim} in the beam.

- On the other hand, the observed source should be far enough from its neighbours to be properly separated; this is the basis of the "source density criterion", which is derived from a completeness criterion and evaluates the density of the sources above the detection threshold S_{lim}, such that only a small fraction of the sources are missed because they cannot be separated from the nearest neighbour.

Generally, we should compare the confusion noise derived from both criteria, in order not to underestimate it artificially. The confusion noise, σ_c, and confusion limit, S_{lim} are defined as follows:

$$\sigma_c^2 = \int f^2(\theta,\phi)d\theta d\phi \int_0^{S\_lim} S^2(dN/dS)dS$$

where $f(\theta,\phi)$ is the instrumental 2D beam profile, that can be approximated by a Gaussian profile with the same FWHM as the expected PSF, or by an Airy function, S is the source flux density (in Jy) and dN/dS is the differential source number counts (in $Jy^{-1}sr^{-1}$).

Then, the total noise is computed by adding in quadrature the different noise contributions, in this case the photon (and instrumental) noise and the confusion noise, i.e. $\sigma_{total} = (\sigma_p^2 + \sigma_c^2)^{1/2}$

The photometric criterion is defined by choosing the S/N ratio q_{phot} between the faintest source (of flux S_{lim} and the noise σ_c due to fluctuations from beam to beam caused by sources fainter than S_{lim}, as given by the implicit equation:

$$q_{phot} = S_{lim} / \sigma_c(S_{lim})$$

q is usually chosen between 3 and 5, depending on the specific objectives.

The source density criterion is defined by setting the minimum degree of completeness of the detection of sources above the limiting flux S_{lim}, which is driven by the fraction of sources lost in the detection process due to a nearest neighbour source with flux above S_{lim} too close to be separated given an instrumental beam size. For a given Poissonian source density N(>S), the probability P of finding a nearest neighbour with $S \geq S_{lim}$ at a distance closer than the minimum angular separation θ_{min} is given by:

$$P(< \theta_{min}) = 1 - \exp(-\pi N\theta_{min}^2)$$

An acceptable probability limit is $P_{max} = 0.1$. The minimum distance is usually parameterised using the FWHM of the beam profile $\theta_{min} = k\theta_{FWHM}$, and $0.8 \leq k \leq 1$. Fixing the probability we obtain the corresponding "source density criterion" limiting density of sources:

$$N_{SDC} = -\ln(1-P(<\theta_{min})) / \pi Nk^2\theta_{FWHM}^2$$

The instrumental beam area, is given by $\Omega \sim 1.14\theta^2_{FWHM}$. Therefore, for P = 0.1 and k = 0.8, the density is 1/16.7 sources/beam. The limiting source flux, S_{SDC} is thus determined by using existing number counts results and a suitable model for infrared galaxy evolution extrapolating the data to the appropriate wavelengths and (faint) flux levels. The confusion noise, σ_{SDC} is computed using the same relation as for the photometric criterion, as the S/N ratio $q_{SDC} = S_{SDC} / \sigma_{SDC}$.

The Herschel confusion noise levels due to extragalactic sources in the different instruments/bands are being computed from deep maps on 'blank' fields. The preliminary results ([RD11] and [RD12]) are shown in Table 4.1 along with the values predicted by by Lagache et al. 2003 ([RD8]) using number counts derived from a phenomenological model based on template spectra of starburst and normal galaxies, and on the local infrared luminosity function. This model has been found to be in very good overall agreement with ISOCAM at 15 μm, IRAS at 60 and 170 μm and SCUBA at 850 μm (see references within [RD8]). Confusion level predictions for Herschel/PACS have been also computed by Dole et al. 2004 ([RD9]) based on recent Spitzer/MIPS number counts from Papovich et al. 2004 ([RD10]) shown in Figure 4.4. They obtain $S_{SDC}(70 \ \mu m) = 0.16$ mJy and $S_{SDC}(160 \ \mu m) = 10.0$ mJy. Please refer to the specific instruments' Observers' Manual for an up-to-date information regarding confusion noise.

Table 4.1. PACS and SPIRE measured confusion noise, compared to predictions computed according to photometric and source density criteria. From [RD9]. I.

	$\sigma_{observed}$ (mJy)		σ (mJy)
PACS 70 μm	N/A	q_{phot} = 5.0	2.26×10^{-3}
		q_{SDC} = 8.9	1.42×10^{-2}
PACS 100 μm	0.27	q_{phot} = 5.0	1.98×10^{-2}
		q_{SDC} = 8.7	1.02×10^{-1}
PACS 160 μm	0.92	q_{phot} = 5.0	3.97×10^{-1}
		q_{SDC} = 7.13	9.93×10^{-1}
SPIRE 250 μm	5.8	q_{phot} = 5.0	2.51
		q_{SDC} = 5.2	2.70
SPIRE 360 μm	6.3	q_{phot} = 5.0	4.4
		q_{SDC} = 3.6	3.52
SPIRE 550 μm	6.8	q_{phot} = 5.0	3.69
		q_{SDC} = 2.5	3.18
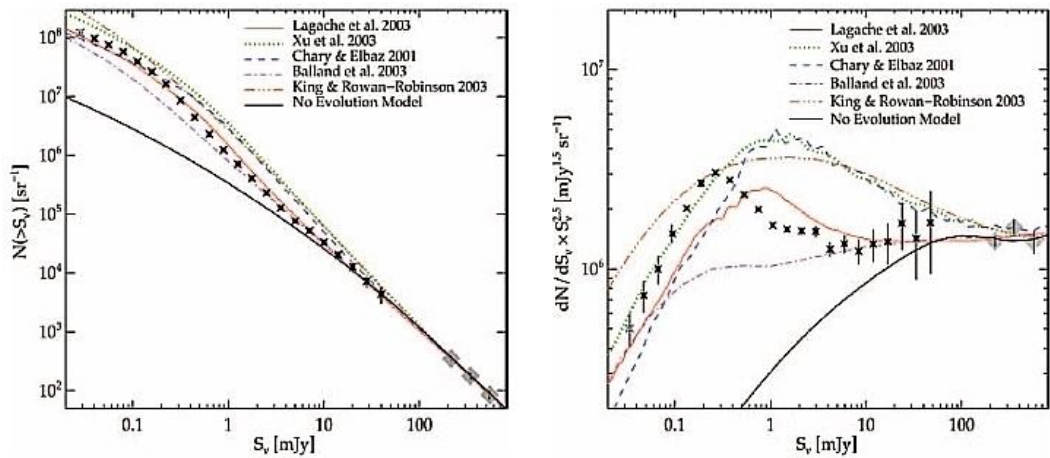
Figure 4.4. Cumulative (left) and differential (right) 24 μm number counts from [RD10]. The differential counts have been normalised to an Euclidean slope, $dN/dS_\nu \sim S_\nu^{-2.5}$. The curves show predictions from different recent models, including that from Lagache et al. 2003.

4.4. Straylight

The Herschel design has been carried out including the instrument optical layout. This approach allows the level of straylight that originates from the various sources at detector level to be provided directly. Therefore, the straylight requirements are given directly as the straylight reaching the detector. The following apply over the full operational wavelength range:

- **Scattered light from sources outside the telescope FoV**: Taking into account the worst possible combination of the positions of the Moon and the Earth w.r.t. the line of sight (LoS) of the telescope, the extreme values are:

  - Sun-S/C-Earth angle of 37°

  - Sun-S/C-Moon angle of 47°

  - Sun-S/C-LoS angle of 60°.8 to 119°.2 (in the S/C XZ plane)

  - Maximum roll angle of ±1°
  The straylight will be < 1% of background radiation induced by the self-emission of the telescope.

- **Sources inside the FoV**: over the entire FoV at angular distances ≥ 3 arcmin from the peak of the point-spread-function (PSF), the straylight shall be $< 1 \times 10^{-4}$ of PSF peak irradiance (in addition to level given by diffraction).

- **Self-emission**: The straylight level, received at the defined detector element location of the PLM/FPU straylight model by self emission (with "cold" stops in front of PACS and SPIRE instrument detectors), excluding the self emission of the telescope reflectors alone (but including any other contributor, notably the M2 hexapod), shall be < 10% of the background induced by self-emission of the telescope reflectors.

According to current straylight analysis for the orbit configuration of Herschel (see [RD3]), for sources outside the FoV, the straylight radiation is within specification, except for small locations on the sky, where radiation reflected from rectangular hexapod structures can enter the instruments directly. These small locations exist primarily for the Moon. Only two minor paths were found which could be applicable also to the Earth. For the worst-case locations of the Moon, the specification is exceeded by a factor 16.4.

For sources inside the FoV, the requirement is met by a wide margin.

Finally, for thermal self-emission, the requirement is not met. Actual values (expressed as a fraction of the background induced by self-emission of the telescope reflectors) are:

- 30% for PACS and 19% for SPIRE (pessimistic case)

- 12% for PACS and 8% for SPIRE (optimistic case)

The HSC has created a dedicated working group to study the straylight effffects on the Herschel operations. New models based on the 'as built' optical system have been prepared and observations designed to verify the models. Figure 4.5 .
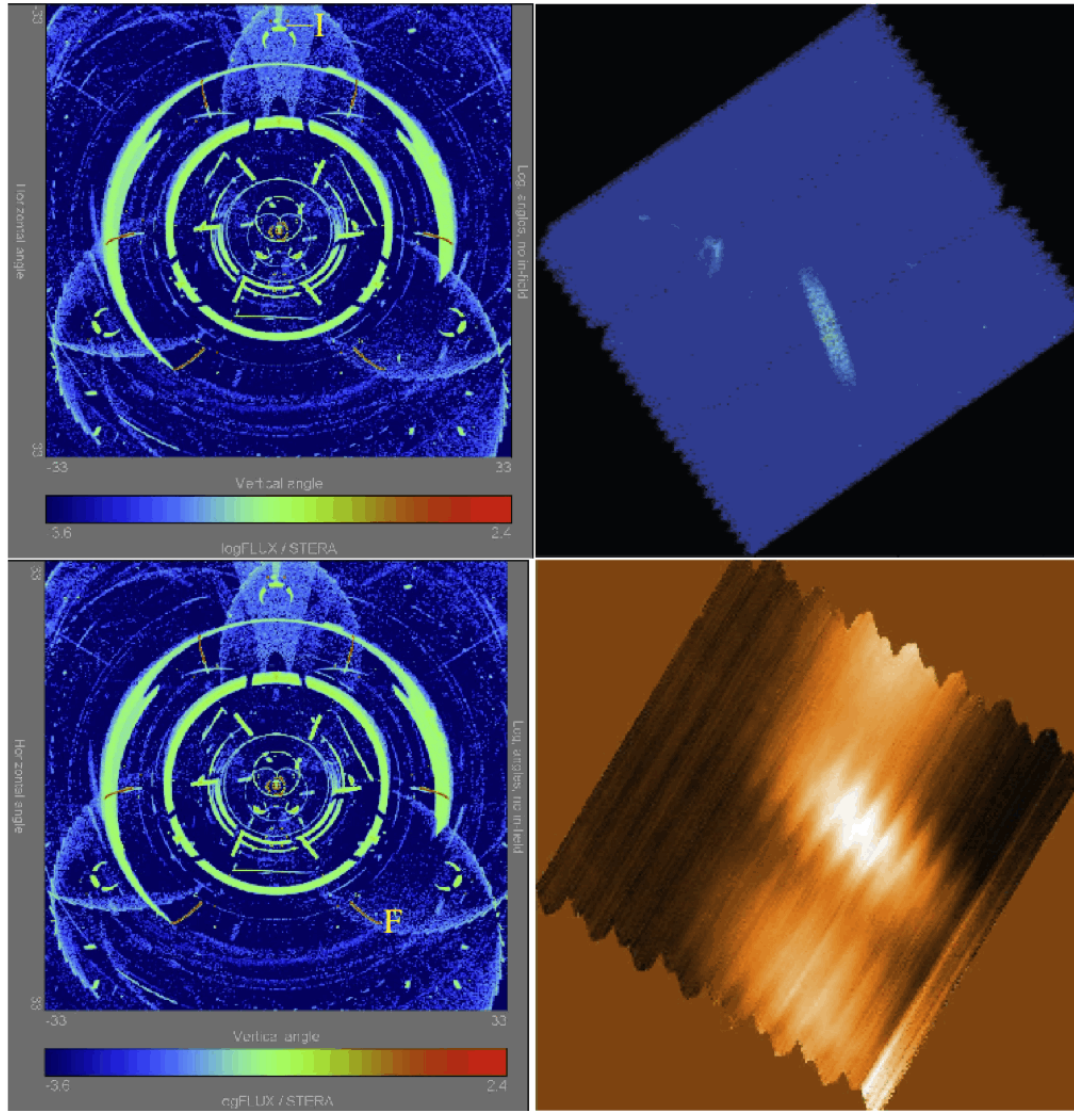
Figure 4.5. Comparison of straylight optical models produced by M. Ferlet (priv. comm.) and observational results. In the top row, a Herschel observation has been planned with Jupiter in position 'I', while in the bottom row the Moon has been placed in position 'F'. In both cases, there is a very good agreement between the model prediction and the straylight results.

Chapter 5. Ground Segment

5.1. Ground Segment Overview

The operations of the Herschel Space Observatory are conducted in a decentralised manner. As can be seen in Figure 5.1, the Ground Segment comprises the following elements:

- A Herschel Science Centre (HSC), provided by ESA, located at ESAC, Madrid. The HSC, supported by the NASA Herschel Science Center (NHSC), located at IPAC, acts as the point of interface to the science community and the outside world in general. The HSC is supported by the Herschel Science Team, for the maximisation of the scientific return of the mission, and by the Herschel Observing Time Allocation Committee (HOTAC) for the selection of observing proposals.

- Three dedicated Instrument Control Centres (ICCs), one for each instrument, provided by the respective PI. Each ICC is responsible for enabling the operation and for the calibration of its instrument.

- A Mission Operations Centre (MOC), provided by ESA, located at ESOC, Darmstadt, which is responsible for the execution of all in-orbit operations.
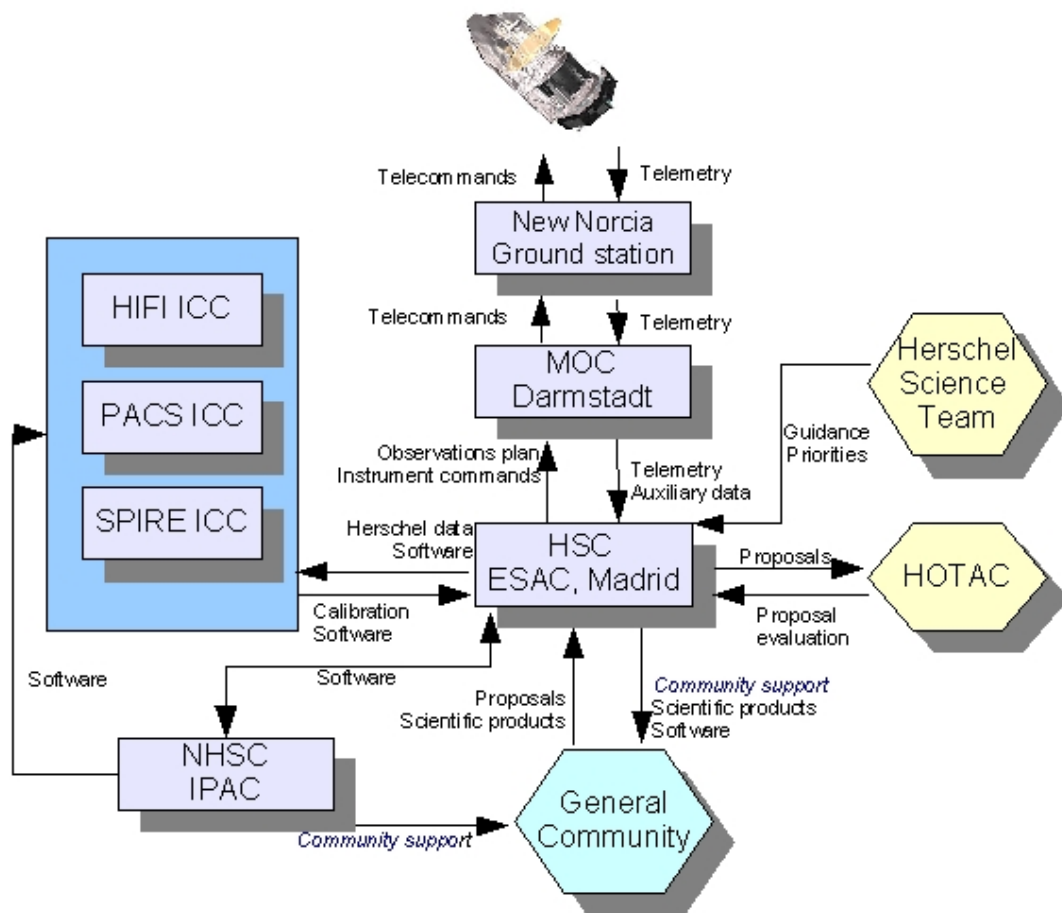

Figure 5.1. Herschel Space Observatory Ground Segment

5.2. From proposal to observations

The Herschel Science Centre provides the information required for the submission of proposals in the Herschel Space Observatory Web site (http://herschel.esac.esa.int), in collaboration with the ICCs. Astronomers are requested to register to access observatory services, which include the capability to submit proposals, access to the Helpdesk and retrieval of observational data from the Herschel Science Archive.

Proposal entering and submission is done through the HSpot tool (see Section 6.1), the Herschel Observation Planning Software. A scientific proposal contains at least one AOR, or Astronomical Observation Request. Each AOR is based on an AOT, or Astronomical Observation Template, which is a pre-defined observing mode, characterised by an instrument configuration and way of operation that have been optimised for the execution of a particular type of observation (see Chapter 6). An AOR is generated when the proposer provides the parameters required for the selected AOT, and is equivalent to the term "observation" used in this document.

A proposal submitted through HSpot is stored in the Herschel Space Observatory database. The proposer, and co-proposers selected by the principal investigator, are allowed to retrieve, modify and upload their proposal(s) until the closing date of the AO. At that time, the database is closed to HSpot, and the HSC distributes the stored proposals to the HOTAC panels. Proposers can check the status of their proposal(s) in relation to the HOTAC review in the Proposal status Web page (http://herschel.esac.esa.int). During the review process, the HSC provides support to the HOTAC and, on request, assesses the technical feasibility of the observations. In addition, a systematic technical feasibility assessment is carried out on all accepted proposals.

The period of proposal submission before the HOTAC review is called Phase-1. After the HOTAC review results are public, proposal submission Phase-2 starts. In this period, observers are allowed to refine their accepted proposals, modify them following the HOTAC guidelines, and use updated AOTs and the latest available observatory knowledge. Please see the "Herschel Space Observatory Call for Proposals: Policies and Procedures" document for a definition of proposal submission Phase-1 and Phase-2, and for the policies on proposal modifications. The end of proposal submission Phase-2 results in a consolidated database of accepted proposals and its corresponding AORs.

5.3. Calibration observations

The calibration and cross-calibration of the Herschel instruments is the responsibility of the observatory, in particular of the ICCs and the HSC. The pointing calibration is the responsibility of the HSC and the MOC. Therefore, the preparation and scheduling of calibration observations is an exclusive duty of these groups. The calibration data required for the reduction and analysis of the Herschel observations will be provided to the astronomer in the form of products in the Herschel Science Archive, and is integrated in the Data Processing software.

Calibration and engineering observations were the main components of the schedule during the Commissioning and Performance Verification phases. Their aim was to achieve the necessary understanding of the instruments and spacecraft, and attain the required calibration and pointing accuracies to ensure a proper execution and data reduction of the science observations during the Science Demonstration and Routine phases. In the routine phase, it is expected that up to 15% of the available observatory time will be used for calibration. Calibration observations may be based on non-AOT observing modes defined by the instrument specialists at the ICCs and HSC, but in general they will be defined using the AOTs available to the community for science observations. Calibration observations are in principle public. However, if a calibration observation is a duplicate of a scientific observation (see the "Herschel Space Observatory Call for Proposals: Policies and Procedures" document for a definition of "duplication"), the corresponding proprietary rights will apply.

Chapter 6. Observing with Herschel

Herschel is an observatory mission. Thus, as in ground-based telescopes, the astronomer who is requesting the observations must provide all the information necessary to carry them out. These instructions are known as an "Astronomical Observation Request" (AOR), which is made using a standard Astronomical Observing Template (AOT) (see Section 6.3). This information is then converted into spacecraft and instrument commands that are uplinked to the spacecraft to execute the observations. An additional complication with Herschel is that communication with the satellite is normally limited to 3 hours each day, so that all the commands to carry out observations must be uplinked at least 48 hours in advance of the observations being carried out and must be executed autonomously. This means that far more detail must be defined by the observer than for observations in a normal groundbased telescope. The system is designed to make the highly complex process of defining observations as simple as possible for the observer. The following section describes this process.

6.1. Introduction to HSpot

The astronomer's interface with Herschel is an observation planning program called HSpot. HSpot allows the astronomer to define targets and observations, to calculate the time required and likely s/n and to submit a proposal with the requested observations. At any stage of this process the work in progress can be saved and recovered later. HSpot has been adapted from the original Spitzer Space Observatory SPOT program and thus will be familiar to Spitzer users. The part of HSpot directly adapted from SPOT is known as the "Spot Core" of the program and is maintained by IPAC (about 80%), while the HSC maintains the layer of Herschel-specific functionality (about 20%); HSpot incorporates a total of more than 30 man-years of work between the two centres.

HSpot can be downloaded from the [Herschel Science Centre web page](#) at the url:

[ftp://ftp.sciops.esa.int/pub/hspot/HSpot_download.html](#)

Alternatively, select the "Tools" option from the left hand menu of the [Herschel Science Centre web page](#) and "HSpot download" in the tools page.

HSpot is eminently user-friendly and simple to use and has many functionalities that are of interest even to non-infrared astronomers. New users can generally familiarise themselves with the main functions in an hour or so of simply playing with the program.

6.1.1. Keeping HSpot up to date

HSpot is updated regularly. For the 2011 Guaranteed Time Call, a completely new and revised version will be released (HSpot 5.3), including literally many minor and major updates since the previous Open Time Call and also numerous updates to the underlying Spot Core. Version 6.0 is projected to be released for the OT2 Open Time Call. Occasionally, unexpected issues come to light requiring a new update of HSpot, in which case a new release is made that will be downloaded automatically from the HSC. For each new Call and at key points between Calls a new version of HSpot will be made available with any necessary updates: the default is that HSpot will download these updates automatically and offer them to you. It is strongly recommended that you do not change this option, as it may lead to submitting or revising your AORs against a wrong HSpot version, or to having incorrect time estimates for your AORs (in extreme cases it could even lead to your proposal being rejected automatically).

Each time that you open HSpot, it will connect to the HSC server and check to see if a new version is available. If one is found, you will be offered a choice of closing HSpot immediately and reopening it with the new version, of waiting to install the new version, or of refusing the update (in which case automatic updates are disabled in the future). You are strongly advised to accept the update immediately; normally it will be installed and operational in under a minute.

Similarly, occasionally a new time estimator version may be announced when you open HSpot; the time estimator links your AORs to the latest instrument control software that sets the parameters for each observation. Normally time estimator changes will not affect the duration of observations now,

but will effect essential parameters in the set-up of observations. When the time estimator version is updated, the time estimate for your previously prepared AORs will be shown in red; it is essential to submit your proposal against the latest time estimator version - to do this, just re-run time estimation before submission so that all the time estimates are shown in black font.

6.1.2. Will HSpot run on my computer?

HSpot has been developed to run on the three main operating systems currently in use: Unix/Linux, Windows and Mac. The development work has been carried out on Solaris and ported to these operating systems and the system has been extensively tested. We thus believe that HSpot should run reliably on all the principal operating systems available to users. For each operating system certain common platforms are supported. Users are strongly urged to use these standard combinations of operating system and platform, as no guarantee can be offered that HSpot will run correctly on other combinations and no guarantee can be made of support for other platforms. Similarly, users will understand that, for example, the Windows version of HSpot has been extensively tested on Windows XP and Vista.

To date, minimal testing has been done on Windows Seven, but we do not anticipate problems with its use, have not been notified of any issues and do know that the installer works correctly on this Operating System, although it has not yet been officially supported. Acceptance testing will be carried out at HSC as standard on Windows 7 from Spring 2011 onwards.

Detailed information on the operating systems and platforms supported can be found in the HSpot manual. HSpot runs under Java and users are strongly advised to ensure that all updates and patches of their operating system are installed.

Warning
HSpot will only run on Java 1.6 and later. Older Mac machines that do not support a Dual Core and 64-bit architecture will not install HSpot. Due to the way that Mac handles Java, Mac users have occasionally experienced minor problems with HSpot that Windows, Solaris and Linux users have not.

6.1.3. Proposal presentation

Proposal presentation is extremely simple with HSpot. Once the observations to be carried out are defined and saved, the proposal can be submitted quickly and easily from the "Tools" menu. A submitted proposal can be retrieved before the deadline for submission and revised as many times as required; this allows you to submit a draft and then update it continuously so that, even in case of disaster (your local hard disk fails, the Internet falls over just before the submission deadline, etc), HSC will always have a valid latest version of the proposal. To submit a proposal, apart from the AORs (that is, the source information, instrumental configuration, exposure time, etc. for each object to be observed) the proposer needs a text file with the proposal abstract (maximum 2000 characters, including spaces), which can be read in directly, a PDF file of the scientific justification (limited to a maximum of 5Mbt and prepared with the HerschelFORM PDFLatex package that is available on the Herschel Science Centre webpage) and to give basic information such as the proposal title, list of co-Is and the observing call that the proposal is responding to.

When a proposal has been submitted HSpot will confirm that it has been transmitted correctly and, on completion of processing, an e-mail will be received from the HSC Proposal handling System confirming its successful receipt.

Warning
The time taken to generate and transmit the acknowledgement e-mail is a strong function of the system load. When the HSC servers are heavily loaded close to a call closure, the acknowledgement e-mail may take tens of minutes or even, in extreme cases, a few hours to arrive. Until this e-mail is received, you will not be able to retrieve and update the latest version of your proposal. All proposals to arrive are logged with time of submission and the HSC will know that your proposal is in the system, even if you have not yet received the formal acknowledgement.

6.2. Types of target

HSpot deals with two fundamental types of target: fixed targets and solar system objects.

6.2.1. Fixed targets

A fixed target is any object that does not require a differential tracking rate. This can be a star, a galaxy, an AGN, etc. Herschel works with Equatorial J2000 coordinates and only target entry in Equatorial J2000 will be accepted (this is to facilitate checks for duplicate pointings [two or more users requesting similar observations of the same object, or region of the sky], which are extremely complicated if many coordinate systems are used for target entry). If the source is known to NED or SIMBAD these coordinates are used, if not, the user must enter a J2000 R.A. and Dec. On some occasions, for nearby stars, the proper motion of the target may become important; this can be entered in HSpot if necessary, once again, the epoch must be in 2000 coordinates. All fields can be edited after name resolution.

Some types of observation can be very unforgiving with the coordinates. Typical problems found are missing "minus" signs in declination and, where several positions are available for a target, picking an inaccurate one. Many faint infrared sources do not have good coordinates and this must be bourne in mind when planning observations. Particularly HIFI observations in the high freqency bands (Bands 6 and 7) are sensitive to poor pointing accuracy and may be affected if the quality of the position of a fixed target is low.

6.2.2. Moving targets and their treatment

A moving target is a solar system object that requires a differential tracking rate to be programmed. On target entry the user should select the "Moving" tab and resolve the NAIF ID of the target name. The Herschel Observations Planning System will use the NAIF ID to calculate coordinates for the time of observation and to calculate the differential tracking rate required, which should be less than 10 arcsec/minute at the date of observation (this limits the capability of Herschel to see objects passing very close to the Earth, although faster rates up to 30 arcsec/min may be permitted, on a case-by-case basis, if scientifically justified). User entry of target coordinates is not permitted, as any solar system object with a reliable enough orbit to be observed by Herschel will have a NAIF ID.

Around 800 moving targets (satellites, comets, asteroids and TNOs) are in the HSpot database. More than a million have been catalogued, but it is obviously impractical to store all of them in HSpot as most are not observable by Herschel. If you wish to observe a solar system object that is not in HSpot you should send a Helpdesk ticket requesting that it be added. Allow a minimum of two or three working days for it to be included and the ephemeris to be linked to HSpot.

6.2.2.1. What is a NAIF ID?

NAIF is NASA's Navigation and Ancilliary Information Facility. This offers an information system called SPICE for spacecraft navigation. SPICE uses a unique 7 digit identification code for all natural solar system bodies, while spacecraft are identified with a negative integer code. Because of the simplicity for this system of ID codes and given the increasing possibility of confusion of objects (for example, there are both planetary satellites and asteroids named Io, Ganymede and Dione and increasing numbers of asteroids are later found to show cometary activity and may receive multiple designations), it is increasingly used for telescope scheduling. A short summary of NAIF IDs is given in the relevant secion of the HSpot Users' Manual on the Standard Ephemeris for moving target entry.

6.2.2.2. Solar system object ephemeris accuracy

When a Solar System Object has a well-controlled orbit of high accuracy (for a periodic comet this means two returns for which a successful linkage has been made, for a asteroid or minor body it usually means observations at a minimum of 6 or 7 oppositions, apart from Earth-crossing objects for which the criterion is typically 3) it will receive a number from the Minor Planet Center. A numbered comet has a designation such as 190P/Name, while an asteroid receives just a number. An unnumbered asteroid has a NAIF ID starting with a 3. Objects with such a designation have a relatively low accuracy ephemeris that may be considerably in error when extrapolated to the future. As

an example, even an object with three oppositions may have a position that has a 3-sigma error of more than 60 arcseconds when extrapolated 5 years into the future. If the spread of observations is unfavourable, or there are few astrometric observations, it may not even be possible to obtain a good ephemeris extrapolation with a 3-opposition orbit. With 4 oppositions the 3-sigma error in the extrapolated position may still be greater than 20 arcseconds over 4 years. This means that faint objects that have not been observed recently may be difficult to locate and identify with Herschel and thus are high-risk observations. It also means that an object may not be centred on the detector so the resultant data quality may be defficient.

6.2.2.3. What accuracy of ephemeris is required?

Three problems are present when there is uncertainty in the ephemeris. In approximate order of increasing importance these are:

- Possible errors in the required tracking rate.

  In general the tracking errors should be kept below 1 arcsecond during the observation.

- Difficulties with photometry

  To be carried out successfully, the target must be centred in the array to within a certain level of accuracy. If this is not achieved, photometry may be difficult or impossible to obtain.

- Problems with target identification

  Not all Solar System Objects have suitably accurate ephemerids; occasionally there may be errors of tens of seconds of arc, minutes or even, for a few objects, degrees in the ephemeris position. In the HSpot Users' Manual a list of solar system objects included HSpot is given in which flags objects with deficient ephemerides at the time of writing. Always check to see if a better orbit is now available.

In detail, the issues that users may find when ephemeris information is uncertain are:

- Tracking

  In general this should not be a problem with distant objects, it may become a serious problem with more nearby ones, particularly Near Earth Objects where it may be difficult to keep the target accurately centred.

- Photometry issues

  For PACS photometry, the source position must be known with high enough precision that it should fall within a bolometer matrix of 52x52 arcseconds. In practical terms this means that the following criteria of positional accuracy should be fulfilled.

  -- For aperture photometry: 15 arcseconds.

  -- For PSF fitting: < 10 arcseconds

  For SPIRE the main consideration is that the FWHM of the detectors is 18 arcseconds and the jiggle amplitude 6 arcseconds: if the positional error is greater than the jiggle amplitude there will be light losses.

  For HIFI it should be remembered that the smallest aperture (that of Band 7b) is 13 arcseconds, thus necesitating centering at the arcsecond level to avoid light losses.

- Target identification problems

  For numbered asteroids the ephemeris should be of sufficient precision in almost all cases.

  For unnumbered asteroids and minor bodies it may be essential to take astrometry to refine the

orbit before observations can be attempted with Herschel.

For numbered and ToO comets, recent astrometry may be essential, depending on the case. A numbered comet will almost invariably require post-recovery astrometry to refine the orbit before observation can be attempted. Recently discovered comets with a short orbital arc will also almost invariably require pre-Herschel observation astrometry to refine their ephemeris.

6.2.2.4. Potential problems with moving targets

Standard practice at the HSC is to download the ephemerids for Solar System Objects from the JPL Horizons database every 4 weeks, as our planning is done in cycles of 2 weeks. This means that in an extreme case the ephemeris information for an object may be as much as 2 months out of date. Normally this does not matter, as the errors will be too small to be significant. If the orbit is well-defined, the error is usually almost entirely in the direction of motion, so a Solar System Object (asteroid or comet) will reach a given point in its orbit slightly advanced or slightly delayed with respect to the ephemeris prediction. For Main Belt Asteroids (MBAs) advances or delays of 10 or more minutes are not unknown; at a typical distance of observation of 2AU this translates into a very small error in the actual observed position on the sky.

For objects that come closer to the Earth and move more rapidly, the error in the ephemeris may be much larger, even though the absolute precision of knowledge the object's position may be an order of magnitude better than for an MBA if it is intensely observed. In one recent case an SSO was found to be more than 20 arcseconds different from the ephemeris prediction published only 2 months previously. For Near Earth Asteroids an ephemeris may become effectively completely unusable in a week or less. Observers should be aware that it is their responsibility to ensure that there is sufficient knowledge of the ephemeris of a target for effective scheduling and warn HSC to take the necessary measures to schedule with the most up-to-date available information.

Normally observations are planned and sent to MOC for uplink to the satellite a minimum of 2 weeks in advance and more usually at least 3 weeks in advance. If knowledge of the object's position is likely to be insufficient at that time, the observer should request - in advance - that the observations be re-planned closer to the date of execution, taking advantage of a better ephemeris.

Warning
It is the responsibility of the observer to warn the HSC, via a Helpdesk ticket, of potential ephemeris problems that might affect scheduling of a Solar System Object. This should be done far enough in advance to be taken into account in the standard planning cycle. However, no re-planning of observations can be contemplated less than 5 days in advance of their execution; if the ephemeris is potentially not robust for 120 hours in advance of execution, the observations must be designed to be robust enough to compensate for any positional errors that may occur. It is not sufficient to assume that the HSC will automatically spot all potential conflicts in advance.

6.3. AOT entry

6.3.1. Using AOTs

An AOT is an "Astronomical Observation Template". This will be familiar to users of ISO and Spitzer. An AOT is a standard observing mode with an instrument that can be translated into instructions for the spacecraft to carry out the observations autonomously. Herschel will observe autonomously between DTCPs, so each observation must be carried out in a standard way that the spacecraft can understand. Thus, for each of the instruments only pre-defined types of observations can be carried out. The astronomer produces an AOR (Astronomical Observing Request) by taking an AOT and customising it for the required observations.

Following the experience of ISO, the number of AOTs has been deliberately restricted to allow observers as many options as possible, without requiring an unwieldy number of observing modes to be calibrated.

The first stage in AOR entry is to define the target. If it is a known object its name can be resolved

with SIMBAD or with NED or, for a solar system target, as a NAIF ID. For unknown names (e.g. start points for scans), J2000 coordinates must be supplied by the observer. After defining the object, the observer should check that it is observable by Herschel by calculating its visibility windows. Bear in mind that when you define your observation further you may end up limiting its visibility to part of the target's unrestricted observing window.

Once the target is defined the observer must then select the required instrument and AOT to be used. Nine basic observing modes are supported: for HIFI, single point (point source spectrophotometry), mapping and spectral scans; for PACS, photometry, line spectroscopy and range spectroscopy; for SPIRE, SPIRE photometer and spectrometer; and the SPIRE PACS Parallel Mode. Each of these modes is further subdivided, HIFI, for example, offers a choice of fourteen different mixer bands. PACS photometry allows three variants including the mini-scan maps which have replaced point-source photometry and chopped raster maps. SPIRE Spectrometer offers point source and raster maps, three choices of image sampling, and four choices of spectral resolution, etc. HSpot will guide you through this process of definition with a series of pull-down menus and pop-up windows.

> **Warning**
> Although Point Source photometry may be selected through HSpot, it is no longer offered as a science mode and remains only for specialist calibration applications. Point source photometry should be carried out using mini-scan maps only, which are far more sensitive for the same integration time.

For each observation there is a basic minimum unit of observing time required; the observer need only specify how many repetitions of this unit time are required -- obviously greater sensitivity is obtained through more repetitions (four integrations will give twice the sensitivity of a single one - although for SPIRE, once you reach the confusion limit, you cannot attain better sensitivity however long you integrate), but the observation takes longer. At any time the "Observation Est..." (Observation Estimate) button can be pressed and HSpot will give an estimate of the total time that the observation will take, including the overheads involved, with a break-down of information about the observation. If the total length of the observation exceeds the maximum permitted, HSpot will give a warning that the observation duration is out of limits.

> **Note**
> The one exception that allows the confusion noise to be beaten is for moving targets - Solar System Objects (SSOs). By using an off position we can subtract out the background and thus eliminate almost completely the confusion noise. For a sufficiently long exposure of an SSO, the target will move sufficiently during the exposure that end of the exposure can act as the "off" for the start; this has proved very effective for detecting extremely faint SSOs down to the limiting sensitivity of the telescope.

The observer can vary the parameters of the observation (more or fewer repetitions, nodding on or off, larger or smaller chopper throw, a wider or narrower range of wavelengths or length of scan, etc.) and see how the time estimate varies. Once an acceptable combination of parameters has been found the observer accepts the parameters that are defined to fix the AOR; this AOR can however be modified later, if necessary.

When a proposal is submitted, HSpot takes the currently defined list of AORs and links them to the proposal. It is thus essential to ensure that the correct AOTs and AORs are defined and that the source visibility and observing time are correct for each target.

6.3.2. Full and limited visibility

The star tracker (the telescope's autoguider) is pointing in the opposite direction in the sky to the telescope. So, when a target is close to opposition, the star tracker is pointing sunwards. It has been found that when the solar elongation of a target is greater than 110 degrees direct sunlight falls on the star tracker baffle, leading to heating and thermal distortion of the star tracker unit. This distortion causes progressive guiding errors that affect not just the current observation, but also all posterior observations until the star tracker has cooled again.

After in-flight study it has become obvious that data quality may be compromised by long exposures in the region of solar elongation from 110 to 119.2 degrees (the "hot zone"). HSpot will show obser-

vations in this range as having "limited visibility". In practice this means that according to the degree of incursion into the hot zone increasing strict limits will be put on scheduling. As a rule of thumb, observations longer than 1 hour will not be scheduled in the hot zone. Observations in the hot zone will be scheduled at the end of the observing day unless that are very short incursions that go only slightly into the hot zone: this allows the star tracker to cool during the Daily Telecommunications Period. Only short and extremely urgent observations will be scheduled at more than 115 degrees elongation. Time-critical observations longer than an hour that only impinge slightly into the hot zone (e.g. SSOs) will be treated on a case-by-case basis.

6.4. Constraints on observations

HSpot allows the observer to define many different kinds of constraints on observations. This may be to observe an object at a certain time, to carry out observations in a certain sequence, or with a certain detector orientation, or to repeat observations at a certain interval. However, observers should be wary of overconstraining their observations and of defining constraints that are not strictly necessary, as each constraint that is added makes an observation more difficult to schedule.

Warning
Overconstrained observations may be impossible to schedule.

Tip
When you add a constraint, you should use the "AOR Visibility" button (double click on the AOR to bring up the pop-up with the button) to check that the AOR visibility with the constraint is as you expect. This button looks at the AOR that you have defined and includes all the factors that may limit its visibility (map size, orientation constraints, avoidance angles, etc.) and gives you the effective visibility of the observation.

6.4.1. Chopper avoidance angles

In all chopped observations there is a certain danger that a nearby bright source could lie in the chop position, which is at 90 degrees to the position angle reported by HSpot. HSpot allows chopper avoidance angles to be defined. If, even when the chopper throw is changed, it is impossible to avoid a nearby bright object then defining a chopper avoidance angle should be considered. A chopper avoidance angle tells the observation planning system that the observation should be scheduled in such a way that the chopper will not chop at this range of angles. This however should be done with great caution as a star that looks bright in a DSS or 2MASS image is unlikely to be bright, even at the shortest Herschel wavelengths. A chopper avoidance angle is only necessary when there is a strong far-IR source present in the reference position.

Over the year the apparent rotation of the sky caused by the Earth's orbit around the Sun makes the position angle of the chopper on the sky change (this is the roll angle of the spacecraft, measured from north through east, using the spacecraft z-axis as reference - the z-axis is perpendicular to the orientation of the long axis of the PACS and SPIRE arrays). In other words, by selecting a chopper angle constraint we are effectively placing a timing constraint on our observations, stating that it may not be made at certain times of year. However, the Position Angle calculated in has a strong ecliptic latitude dependence. For sources in the ecliptic the Position Angle will barely vary with time during a visibility window. For the two observing windows available each year two values differing by exactly 180 degrees will be found (Figure 6.1). In these cases defining a chopper avoidance angle is, at best, irrelevant (as the PA will only vary in a range of a few degrees anyway) and, at worst, catastrophic because it is may make all observations totally impossible, with no part of the visibility window permitted.
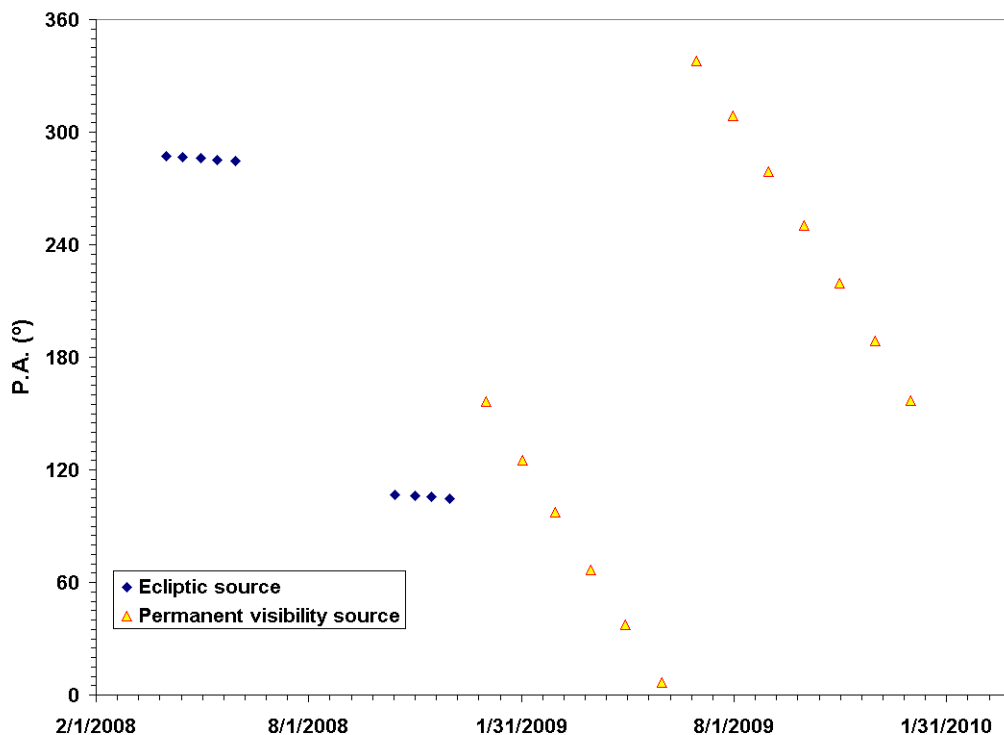
Figure 6.1. Position angle variation for sources on the ecliptic and at the ecliptic pole, in the zone of permanent sky visibility. For sources at intermediate ecliptic latitude the annual range of variation of PA will be between these two extremes. These plots were made originally for a Herschel launch in 2007, but the range and timescale of variation remains unaltered for the actual launch date.

Note

Understanding chopper avoidance angles

HSpot reports the spacecraft roll angle for any particular date of observation. The chop angle will be perpendicular to this angle. If, when you visualise an AOR, you find a bright source in your reference position, you must ADD 90 degrees to the PA in HSpot to avoid a position in the chopper off position. If you have a source in the nod off position you must SUBTRACT 90 degrees to the PA reported in HSpot.

At high ecliptic latitude we have a zone of permanent sky visibility and the PA of the chopper rotates rapidly with time. Here, even a quite wide chopper avoidance angle range may equate to only a relatively small effective restriction on dates. Figure 6.1 shows how the PA changes for a source almost at the ecliptic pole, which is within the permanent sky visibility zone.

At intermediate ecliptic latitudes there will be a break in the visibility windows, although this may be small. When the instrument +Z-axis crosses celestial north there will be a discontinuity in the PA value. Observers should take care of this when defining chopper avoidance angles for sources that are close to +60 degrees ecliptic latitude. A practical example of this is shown for PACS in Figure 6.2 for an object at an ecliptic latitude of 59.5 degrees, close to the point at which there is continuous visibility, but where there is are still two annual visibility windows with a short gap between them. PA=000 degrees is shown (the horizontal position), along with the plotted positions of the PACS imaging detectors are for what was the hypothetical case of a 2007 launch of Herschel, with 2008 March 31st (start of visibility window) PA=127.4 degrees, 2008 June 15th (mid-window) PA=054.6 degrees, 2008 September 10th (end of visibility window) PA=333.7 degrees. The timescale and amplitude of variations does not change for the actual launch date.
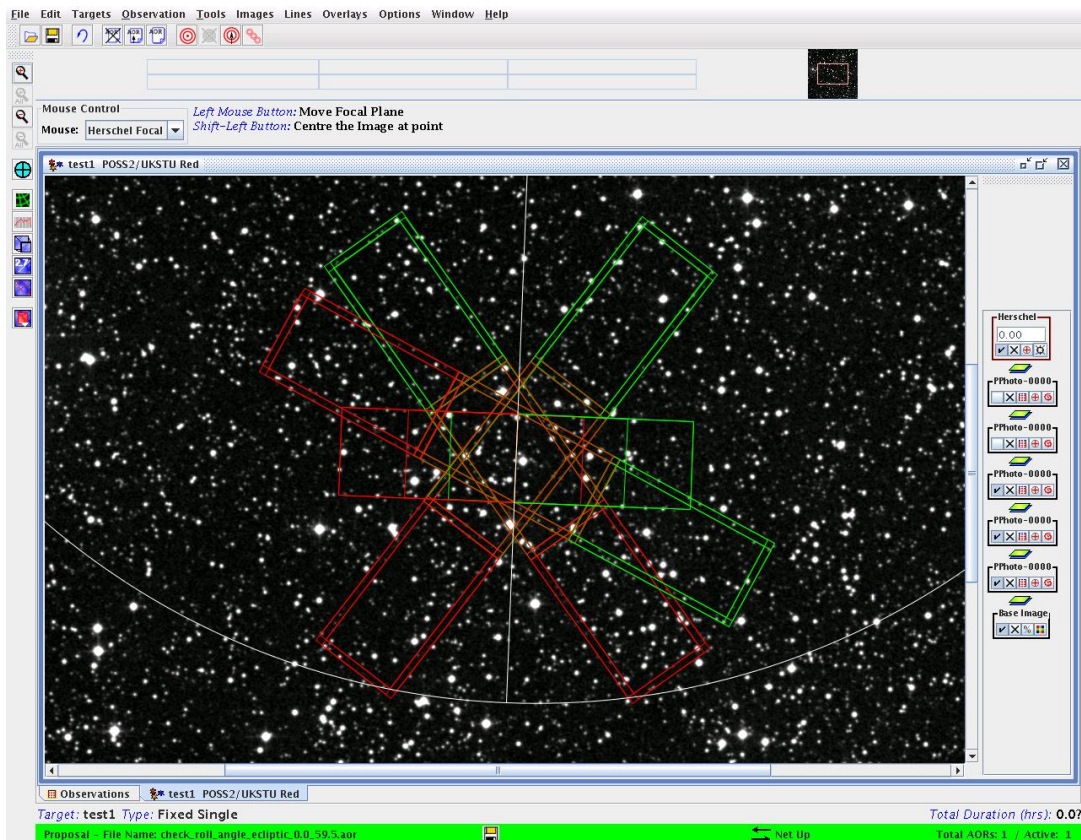
Figure 6.2. An illustrative example. The position angle variation for PACS for an object at an ecliptic latitude of 59.5 degrees, close to the point of permanent visibility. The horizontal position is PA=000 degrees. The plotted positions of the PACS imaging detectors are for a hypothetical case with 2008 March 31st (start of visibility window) PA=127.4 degrees, 2008 June 15th (mid-window) PA=054.6 degrees, 2008 September 10th (end of visibility window) PA=333.7 degrees. The situation is effectively identical for other dates.

> **Warning**
>
> Close to the ecliptic even a small range of chopper avoidance angle may equate to a huge scheduling restriction, potentially making observations impossible to schedule. However, given the very small range of Position Angle change close to the ecliptic, any chopper avoidance angle will either be irrelevant (the PA will never be within the defined avoidance), or catastrophic (the avoidance angle range makes the observation impossible by definition by covering the entire range of PA change).
>
> At high ecliptic latitude it is easier for telescope scheduling to take a chopper avoidance into account.
>
> However, at high ecliptic latitude the chopper PA will often rotate through 360 degrees giving a dephase that must be taken into account when defining a chopper avoidance angle.
>
> In all cases an observer should consider very carefully if defining a chopper avoidance angle is really, genuinely necessary.
>
> All constraints on observations imply an increased observing overhead and thus decreased observing efficiency.

6.4.2. Map orientation constraints

PACS and SPIRE offer the possibility to define a map orientation constraint. In other words, the telescope should scan in a certain direction only, or within a certain range of directions. Further de-

tails of such orientation constraints and their limitations can be found in the relevant instrument manual.

Warning

An map orientation constraint equates to a telescope scheduling restriction and implies that an observation may only be made at a certain, limited range of dates, thus making their execution more problematic. Over-restricting observations may mean that for operational reasons it becomes impossible to carry them out.

6.4.3. Fixed time observations

In certain cases there may be a strong scientific reason for requesting that an observation be carried out at a fixed time. A flag can be put in the AOR defining that the observation be carried out at a set time defined by the astronomer. This obliges the observation planning system to block the observation at this date and time, usually to within a few seconds, although at the cost of putting severe constraints on telescope scheduling, particularly as instruments have to be blocked by days.

A less constraining way of fixing the time is to define a timing window during which the observation should be carried out. A range of dates may be defined during which the observation must be made. This gives the observation planning system more liberty to work around the constraint.

6.4.4. Concatenation of observations

Concatenation or chaining of observations may be defined to oblige the observation planning system to carry out observations together. Concatenation improves planning efficiency by avoiding the need for unnecessary slews, so the observer benefits because no slew overhead is applied to the observation).

Note

The saving may not always be exactly 180 seconds because some set-up is done while the telescope is slewing and so, if a set-up needs to be done for the second, or later observation in a concatenation - for example, an internal calibration - this time will still be charged against the observation.

Concatenation is essential for scan maps, or mini-maps where there is a need to scan in the normal and the crossed direction, to oblige the two scans to be made together and may be convenient in many other cases. This may also be important in the case of a variable object where it is essential that two or more observations are carried out as close to each other in time as possible (an example of such a case might be the need to obtain photometry with PACS at 60-85μm, 85-130μm and 130-210μm, requiring two AORs to be defined that might otherwise be carried out on different days); even for non-variable objects, it is convenient to concatenate observations with PACS at 60-85μm, 85-130μm to avoid slewing away from the object and then back again to take the observation in the second filter, thus adding unnecessary overhead to your observations.

Four methods of chaining of observations are permitted:

- Concatenation of observations

  Two or more AORs for the same target are linked together (concatenated). These must use the same instrument and the same observation type (i.e. you cannot combine PACS and HIFI spectroscopy in a single chain, nor can you combine SPIRE photometry and spectroscopy in a single chain, nor SPIRE PACS Parallel Mode with any other PACS, SPIRE or HIFI mode). At present HSpot does not permit observations in different HIFI bands to be chained either. You can mix a SPIRE photometry map and point source photometry, or a PACS Line Spectrum and a Range Spectrum. The mission planning system will treat these observations as a single pointing. If it is important for observations to be carried out together, they should be concatenated.

  Targets must be separated by no more than 1 degree to be chained. When several pointings are included in a chaining, all must be within one degree of the first position to be defined in the chain to be valid. Fixed and moving targets can be chained, although it is the observer's responsibility to ensure that they will be less than 1 degree apart at some point during the mission and

thus that the observation is schedulable.

Note

There is one exception to the 1 degree rule: to maximise the scientific efficiency of PACS un-chopped spectroscopy the reference position for an AOR may be up to 2 degrees away.

As many chains as are required may be defined and as many observations as are required may be put in each chain, but the total observing time requested in each chain must be less than 18 hours.

The great advantage for the observer, apart from ensuring that observations are carried out together, is to avoid the need for a slew between integrations, thus saving a 180 or 600s slew overhead.

- Follow-up observations

This mode is for repeat observations, for example of a variable source. A time between repeat observations can be defined. Chained observations can be cloned so that the entire chain is repeated after a number of hours or days. The chain or sequence can be repeated several times if monitoring is required over a period of time.

Warning

The observer can request that a sequence be carried out with a very exact interval, or within a band of time (e.g. each observation should be within 8 and 12 days of the previous one). The stricter the constraint, the more difficult it will be to accommodate the observations in the observing schedule, to the point that highly constrained observations may be impossible to carry out. There is a regular planning cycle of instruments over each two week period, with instruments available on set days in each period: your constraints should be compatible with this cycle.

- Sequencing

This mode is to carry out observations in a particular order, although not necessarily the same day. This may be necessary when two or more measurements are required and it is essential that one be carried out first to allow the other observations to be reduced when carried out.

- Group within

In this mode observations must be carried out in a certain time frame, but with no constraint as to *when*. An observer can specify that all the observations in the group should be carried out within a maximum of, for example, one month; in this case the observatory planning system will complete all the AORs within a month of carrying out the first one. The observations may be carried out in any order within this time interval.

6.5. Limiting length of observations

6.5.1. Fixed targets

There are a series of fundamental constraints on the length of observations with Herschel. There is an operational constraint that the coolers on PACS and SPIRE must be recycled for approximately 2 hours every 48 hours. However, in practice, the limit is imposed by the need to have a 3-hour daily telecommunications period (DTCP) with the ground station to download data and upload instructions every day and up to 3 hours each day are also reserved for routine calibrations. Thus there is a limit of 18 hours to individual observations with Herschel that is hard-wired into HSpot. Observers who wish to take longer observations than this must split their AOTs into shorter segments. Special care should be taken when requesting observations close to the 18 hour limit that they will remain possible even if there is a slight change in on-board observation or calibration strategy as knowledge of the instruments improves in-flight. Very long AOTs impose strong constraints on mission planning and may be difficult to accommodate in the telescope schedule because they will effectively fill an entire observing day and block it for other observations. However, for a photometric deep integ-

ration on a fixed target, the telescope can only stare at a single point in space for 50000s (13.9 hours) thus, the maximum point source photometry AOR length is significantly shorter than 18 hours.

6.5.2. Moving targets

Moving targets must be dealt with in mission planning in a different way to fixed targets, as the spacecraft must calculate an instantaneous position and track on it, rather than on the stars. This requires the mission planning software to interpolate the position of the object at any moment from the Chebyshev Polynomials that define the target's ephemeris. This process may not be valid for integrations longer than 5 hours on fast-moving targets, hence the tracking accuracy cannot be guaranteed for longer moving target AORs; HSpot places a blanket limit of 5 hours on the duration of a single AOR for solar system objects.

6.6. Observing overheads

Each observation that is made with Herschel implies certain overheads. These are detailed in the time estimation breakdown and are charged against the observation. The onus is thus on the observer to make observations as efficient as possible so that precious observing time and thus irreplaceable helium is not wasted on unnecessary overheads.

6.6.1. Telescope slew time

Herschel takes a certain amount of time to slew between targets. The median slew time has been found in the early phases of routine observing, as expected, to be of the order of three minutes (although this depends critically on the density of targets in the sky, which differs for different instruments), thus all unconstrained observations are charged 180s as observatory overhead for slewing the telescope (for constrained observations a 600s slew overhead is applied - see Section 6.6.4). It is still possible that at a later date the 180s median slew overhead will change as the observing database is filled to its maximum extent and our knowledge of source distribution and the effective distribution of slews on the sky becomes more complete. For concatenated observations on the same target a zero telescope slew overhead is applied.

6.6.2. Scans and rasters

When making maps there are certain overheads implicit in the process.

6.6.2.1. Raster maps

In a raster map the telescope must make a slew, stop and wait for the pointing to be stabilised. Due to the satellite's large moment of inertia the process of acceleration, deceleration and stabilisation adds a significant dead time (of the order of 5s) to the measurement in each position. This value has been optimised in the light of in-flight experience and is now not likely to change further.

6.6.2.2. Scan maps

Scan maps have generally been more efficient and added less overhead to an observation than a raster map, although for a scan map the calculation of the overhead is uses a complex formula because several variables are involved. In this case the overhead is the acceleration at the start of a scan and the deceleration at the end of the scan, which will vary according to the length of the scan itself (for short scan legs the telescope will spend a much larger fraction of the time accelerating and decelerating). The telescope then makes a small slew to the start position for the return scan.

Although mini-scan maps look inefficient due to these overheads compared to point source photometry, in fact they go significantly deeper in the same total time and produce much more accurate photometry.

6.6.3. Internal calibration

Each observation requires an internal calibration against black body sources maintained at rigidly controlled temperature. These measurements are essential to the health and success of all observations and are thus charged against the observation. The calibration time is typically in the range 30-300s according to the AOT used.

If the calibration time is less than the slew overhead, it is not charged to the user as an overhead as the calibration is carried out in its entirety during the slew; when this calibration time excedes the slew overhead that has been applied, the excess is charged as an overhead to the astronomer. Obviously, if two observations are concatenated and no slew is involved, the whole of the calibration block has to be charged against the observation; for this reason there may still be a small overhead on concatenated observations.

6.6.4. Constrained observations

Constrained observations (see Section 6.4) limit the telescope scheduling and limit observing efficiency, producing what are effectively hidden overheads (e.g. the telescope is forced to slew to a point on the sky that would not be picked otherwise, making the scheduling less efficient), thus a flat rate of 600s will be charged on all constrained observations, in addition to other observational overheads.

If a constrained observation is concatenated, the 600s overhead is applied only to the first observation.

For a fuller definition of what constitutes a constrained observation that will be charged a 600s overhead, please see the (Policies and procedures) document.

6.7. Details to take into account in the observation of moving targets

6.7.1. Background and PA variations

For all targets the main components of background are the zodiacal light (at short wavelengths, with only slow angular variations and little granularity) and the Interstellar Medium (ISM) at longer wavelengths (with much greater granularity). For a fixed target the ISM will have a fixed value at any wavelength, being highest for targets in the Galactic Plane and the zodiacal light will vary with ecliptic latitude and solar elongation. For a moving target the ISM background will, logically, vary with time, although these variations will be a function of the object's heliocentric and geocentric distance - for distant planets the time variations will be slower but, as a corrolary is that an object will take longer to escape from a region of bad background. Note that Infrared Cirrus is highly structured and this structure will affect observations of faint targets. For very faint solar system targets, or where high signal-to-noise is essential, a careful examination of the cirrus may be necessary to look for a hole in the background that will allow deeper observations; once a suitable hole is identified, you can put a time constraint on your observations to ensure that they are made against it - HSC Mission Planners will make great efforts to satisfy such requests, when properly justified.

As an example, the following shows how the PA (Figure 6.3) and the estimated background at 80 microns (Figure 6.4) vary through a visibility window for the satellite Triton of Neptune (NAIF ID 801). At this wavelength the zodiacal light dominates and increases as the solar elongation decreases. Note too how the PA barely changes over the duration of an observing window, meaning that the chopper throw is almost fixed in direction with time; this has strong implications for any potentially orientation-constrained observations.
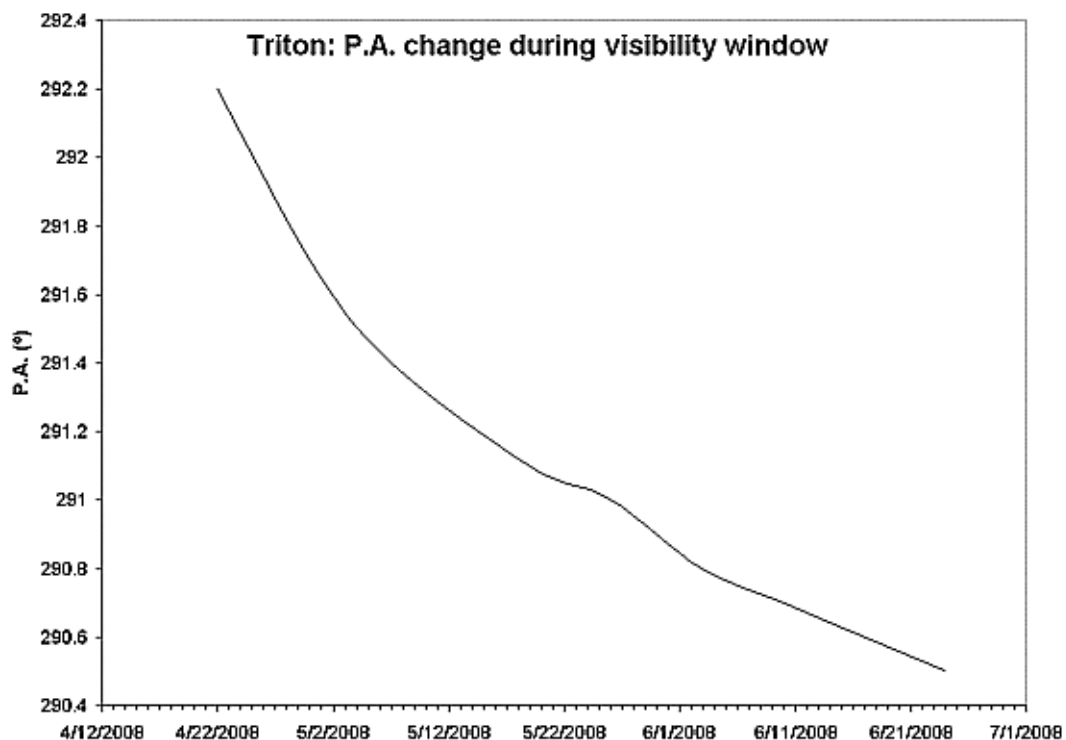
Figure 6.3. PA variation for a typical solar system object: Neptune's satellite Triton. Note how the PA variations over the course of a full observing window amount to less than 2 degrees. This makes it effectively impossible to accomodate map orientation or chopper angle avoidance constraints. Although this example was calculated originally for a Herschel launch in 2007, the amplitude and timescale of variation remains the same for the actual launch date.
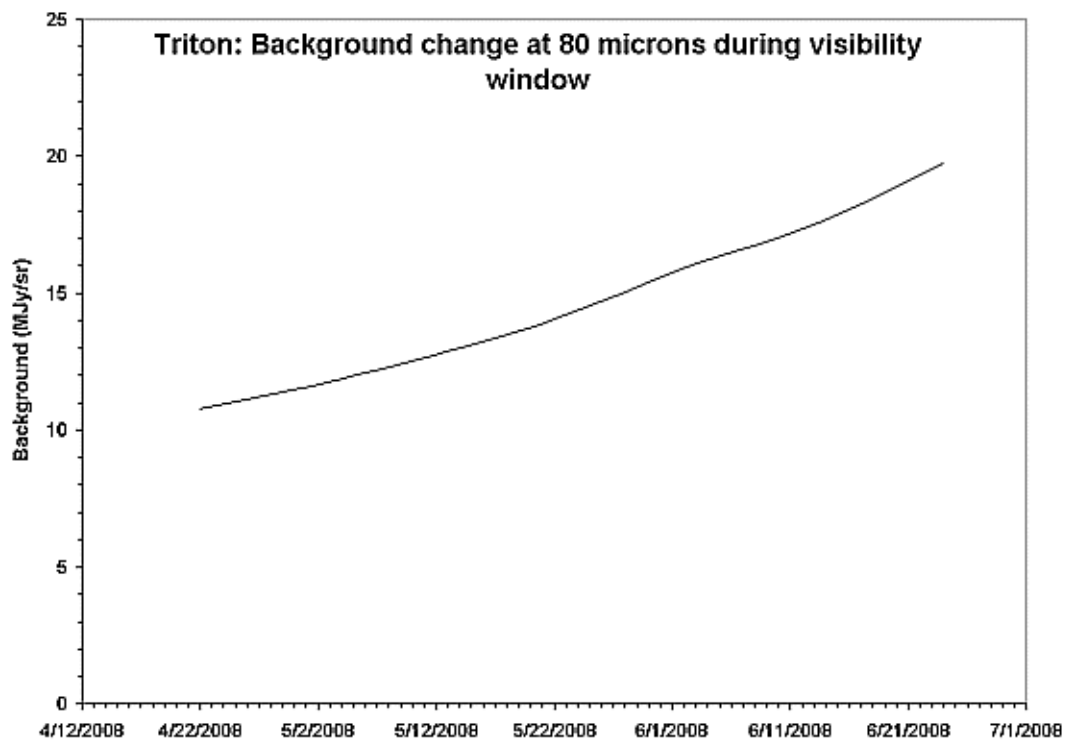

Figure 6.4. The background variation for Triton at 80 microns. The background is dominated at this wavelength by the Zodiacal Light contribution. As the elongation changes over the course of the ob-

serving window the background effectively doubles with time. At longer wavelength the ISM component will also change as the target moves across areas of different background. For objects relatively close to the Sun the ISM component may vary enormously in a comparatively short space of time. Although this example was calculated originally for a Herschel launch in 2007, the amplitude and timescale of variation remains the same for the actual launch date.

6.7.2. Satellite visibility

Note that for satellites of solar system objects HSpot only calculates the visibility window with a solar elongation criterion. It does *not* take into account if the object is genuinely observable by Herschel. It is the astronomer's responsibility to make the necessary checks. Many solar system satellites experience transits and occultations by their parent planet. Similarly, a satellite may not be resolved at the wavelength of observation, or instrument safety constraints may make it impossible to observe a satellite when at less than a certain elongation from the parent planet, or only on one side of the planet. Please contact Helpdesk (http://herschel.esac.esa.int/esupport/) for specific, detailed enquiries about this topic.

> **Warning**
> In photometry mode, instrument safety will not allow Jupiter or Saturn to enter the field of view at any point of the observation.

As an example, the following plots show how the elongation of Io, Jupiter's innermost Galilean satellite (NAIF ID 501), varies from the centre of the disk of Jupiter. In the first plot (Figure 6.5) we see how the elongation varies with time over part of a visibility window. In the area marked in grey the satellite is either in transit, or occulted and thus, by definition unobservable. The second plot (Figure 6.6) shows the offsets in R.A. and Dec. (in arcseconds) over a full observing window. The ellipse marks the approximate size of the disk of Jupiter which suffers a variation of about 10% with time. Note that the entire area of the plot is smaller than the PACS or SPIRE instrument array (see Table 3.1).
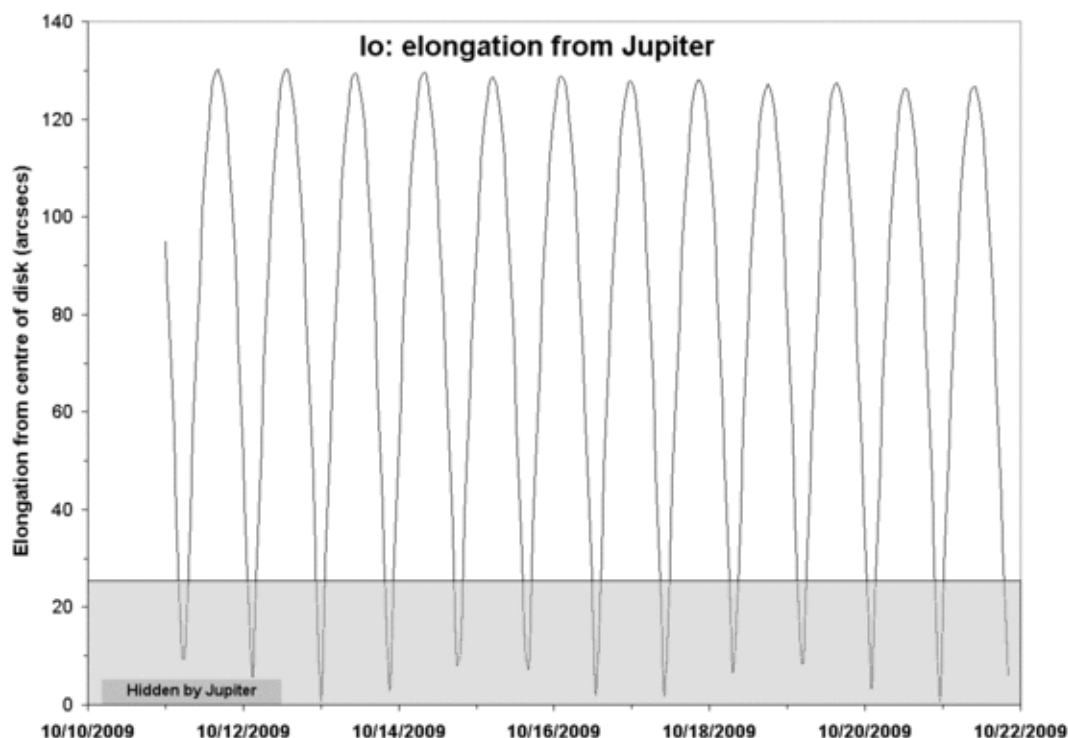

Figure 6.5. The variation of the elongation of Io from the centre of Jupiter with time. The area in grey is the region when Io is either superimposed on the disk of Jupiter (in transit) or behind the disk of Jupiter (occulted). HSpot does not warn the user if visibility of a planetary satellite is limited in this way.
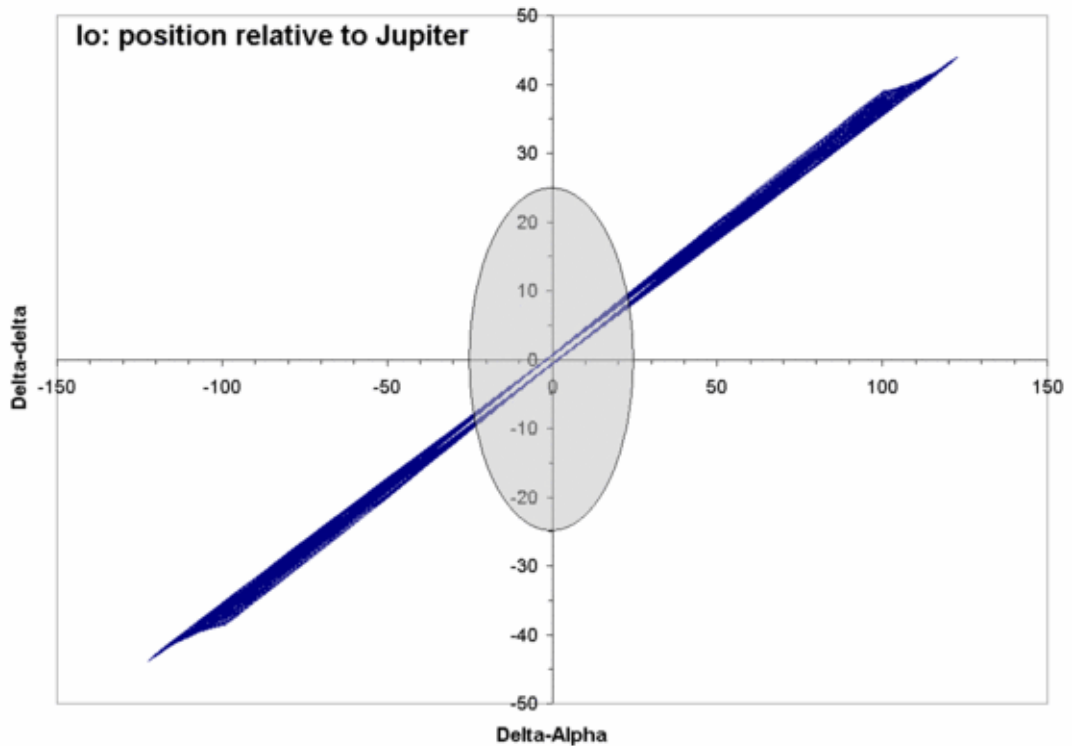
Figure 6.6. The variation in the offset of Io from the centre of Jupiter through an entire visibility window. The grey ellipse represents the approximate mean size of the disk of Jupiter. Note that the entire area of this plot is smaller than the field of view of either PACS or SPIRE. If requesting observations of a planetary satellite the observer should check the visibility of the satellite using the JPL Horizons program at the url: http://ssd.jpl.nasa.gov/horizons.cgi.

Warning
If requesting observations of a planetary satellite the observer should check the visibility of the satellite using the JPL Horizons program at the url: http://ssd.jpl.nasa.gov/horizons.cgi. The ephemeris should be requested specifically for the "Herschel Space Observatory" (site code "500@-486). The observations will almost certainly have to be entered in HSpot with a time constraint save for small, distant satellites.

Chapter 7. Mission Planning and Observation Execution

7.1. Mission planning activities

The observatory schedule is defined by the database of accepted observations. The HSC carries out a careful study of the observation database to define a long-term mission plan that will accommodate all constraints and will maximise the scientific return. The Long Term Mission Planning tool developed at HSC is a very powerful aid to identifying potential future scheduling problems at a very early stage.

Following the agreed long term mission plan, short term observing schedules, together with the corresponding instrument commands, are produced with the Mission Planning System at the HSC, and transferred to the Mission Operations Centre (MOC), at ESOC. MOC adds the satellite commands and produces the final detailed mission timeline that is uplinked to the spacecraft.

The basic time unit for the mission planning is the Operational Day, or OD, defined as the interval of time between the start of two consecutive DTCPs. The DTCP, or Daily TeleCommunication Period, is the time interval when the spacecraft antenna will be pointed to the Earth to receive telecommands and send the recorded data. The duration of an OD will normally be about 24 hours, but depends on the availability and detailed schedule of the New Norcia Ground Station, which is shared with other ESA missions. The operational constraints on the Herschel instruments determine that only observations that use a particular sub-instrument are scheduled in a single OD. For sub-instruments that require cooler re-cycling, only observations with the cooled sub-instrument (e.g. PACS photometer) will be scheduled for the duration of the cooler hold time: two to two and a half consecutive ODs.

Figure 7.1. The default Mission Planning Cycle currently used in routine phase

For a wide range of reasons, from safety to calibration needs, the instrument assignation for each OD is standardised. It consists of the repetition of 28 ODs, i.e. four weeks, during which the instruments follow one another and are used for a different number of consecutive ODs. The standard instrument distribution currently used currently during the routine phase is the one shown in Figure 7.1, which reflects the relative usage of each sub-instrument in the approved proposals so far, and can be revised according to the each instrument observing time preassure at different epochs of the year. This is what we call the "planning cycle", which is the foundation of most of the Ground Segment activities related to Mission Planning. This translates into an additional difficulty to acco-

modate observations that have been defined with timing or grouping constraints shorter than a few weeks.

The use of the different HIFI bands is an additional constraint on the optimisation of observatory time. Very limited HIFI band changes are allowed in a given OD, preferably no more than two. Therefore HIFI observations using different bands and with time constraints shorter than an OD will be very difficult to schedule.

As indicated when requesting the visibility window of an observation with HSpot, it is limited by the so-called "warm" attitudes, i.e. solar aspect angles between -30 and -20 degrees, in which the Sun warms the star trackers and the pointing accuracy may be de-graded (see Section 2.4). For this reason the scheduling of observations within this area is to be avoided where possible. Only in certain, justified occasions, are solar aspect angles between -25 and -30 degrees allowed for less than one hour, at user's own risk and only when it can be demonstrated that it will not affect later observations, while potential scheduling at solar aspect angles from -20 to -25 degrees is decided on a case-by-case basis.

7.2. The execution of the observations

The satellite executes autonomously the mission timeline that has been uplinked during the DTCP. The observational data is stored on board, and downlinked to the New Norcia Ground Station (which is backed-up by the Cebreros Ground Station) during the next DTCP. During this period, which lasts approximately 3 hours, the status of the satellite will be monitored and operational or emergency procedures will be applied, when necessary. In addition, the mission timeline with the commands to be executed during the next OD will be uplinked. This though is a rolling process. In case a DTCP communications linkage with the ground station is missed, the satellite must always have two operational days of observations stored in the onboard solid-state memory. This means that the commands to be executed are always added to the end of the onboard file so that, in the case of a communications failure, there will always be sufficient commands on board to last until the end of the DTCP of the following OD.

It takes several weeks to complete the planning of an OD, from the compilation of all the inputs needed, to the building of the mission timeline by the MOC. Therefore if there are special scheduling requests they should be sent via Helpdesk two months in advance in order to be able to apply them at due time. Due to the complexity of the process, two to three weeks in advance, a draft schedule will be at user's disposal for him to be aware of the imminent execution of his or her observations and to be able to request "last chance" changes in view of availability of recent data. Only in special circumstances will changes be authorised on a shorter timescale.

The downlinked satellite telemetry is transferred from the ground station to the MOC, where it is consolidated and be made available to the HSC. The HSC retrieves the consolidated telemetry and auxiliary data from the MOC automatically, and ingests them in the HSC database, propagating the data to the Instrument Control Centres for each of the three instruments.

Chapter 8. Herschel Data Processing

The scientific analysis of the Herschel observations requires the handling of the Herschel Data Products, which are stored in the Herschel Science Archive. This chapter explains the contents and structure of these products, which include raw and processed data, plus calibration and quality information. It also presents the infrastructure that generates and store them as well as the software provided to the users to analyze them.

8.1. Herschel Data Products

All Herschel telemetry and auxiliary data will be automatically processed at the HSC with the Standard Product Generation software (SPG), to produce the observational data products, stored in the Herschel Science Archive (HSA). The following four levels of Herschel data products are defined:

- Level-0 data product: Raw telemetry data, as measured by the instrument, minimally manipulated and ingested as Data Frames into the mission data base/archive.

- Level-1 data product: Detector readouts calibrated and converted to physical units, in principle instrument and observatory independent. It is expected that level-1 data processing can be performed without human intervention.

- Level- 2 data product: Level-1 data further processed to such a level that scientific analysis can be performed. For optimal results many of the processing steps involved to generate level-2 data may require human interaction, based both on instrument understanding, as well as understanding of the scientific aims of the observation. These data products are at a publishable quality level and should be suitable for Virtual Observatory access.

- Level-3 data product: These are the publishable science products where level-2 data products are used as input. These products are not only from the specific instrument, but are usually combined with theoretical models, other observations, laboratory data, catalogues, etc. Their formats should be Virtual Observatory compatible and these data products should be suitable for Virtual Observatory access.

While the generation of level-0 and level-1 data products will be automatic (although, in some circumstances, some manual intervention may be needed even in Level 1 products), good quality Level-2 products often need a degree of manual intervantion. Level-3 data products can only be generated by interactive processing. It is expected that the degree of human intervention necessary to generate these products will decrease with time, as the knowledge of the instruments' behaviour increases during the mission. This is the same as saying that the quality of the automatically generated products will be progressively enhanced. However, in many cases, it will not be possible to discard interactive processing, especially in the derivation of level-3 data products.

In addition to these observational products, calibration, auxiliary and quality control products will be provided. For more information on the Herschel products, please see the corresponding Instrument Observer Users' Manual. The Herschel Products Definitions document and the Herschel Data Users' Manual document contain detailed descriptions of all Herschel data products.

8.2. Standard Product Generation

The HSC receives, via MOC, the raw telemetry downlinked from the Herschel spacecraft on a daily basis after each DTCP (Daily Telecommunications Period). This raw telemetry set, corresponding to the observations performed in the last 24 hours, the so-called Observational Day, or OD, is ingested in the local database at the HSC and simultaneously propagated also to the main ICCs for quick-look analysis activities. A watchdog setup in the data processing system at HSC monitors the system status after each DTCP ends and automatically decides the readiness of the system to process the data from that OD based on specific criteria being met. When these criteria are met (completion of

telemetry ingestion for that OD in the local database and pointing Product availability in the system for that OD), the watchdog system launches the Standard Product Generation (SPG) pipelines, one for each observation on that OD, in the distributed computing system (GRID) available at the HSC. This distributed system uses several worker nodes containing multiple processors that allow parallel data processing, and is an essential and very powerful infrastructure for the reduction of the large volumes of data generated by Herschel observations daily. During the automatic data processing of one observation, level 0, 1, 2 and quality control products are generated. These data products, together with the inputs used in the processing (auxiliary and calibration data) are 'bundled' together in a so called observation context, a top level container which is the main output of each individual pipeline processing. When the pipeline processing finishes and this top level container is created, it is immediately ingested in the Herschel Science Archive, making it available for retrieval by observers. The Quality control cycle then starts then on the generated pipeline products, following established quality control procedures specific to each observing mode.

The whole process, from observation of an astronomical object to its automatic data reduction using pipelines and ingestion into the science archive, typically takes less than 48 hrs. This is one of the unique features of the Herschel mission in which data products are processed and made available to observers shortly after the data has been taken.

8.3. Quality control

Observation quality control is an important responsibility of the HSC. Its main purpose is to ensure that the observations have been correctly executed, that their observational data meet the established requirements, and that they can be processed error free. It is important to note that the HSC will not assess systematically the scientific validity of individual observations, but will concentrate on their execution and the data processing aspects.

In combination with the SPG processing, the observational data will be run through the Quality Control Pipeline (QCP). An HSC operator will inspect visually all scientific Herschel observations and will proceed according to agreed observatory procedures. For certain types of problems, the operator will request the assistance of the instrument and satellite specialists at the HSC, ICCs or MOC, who will investigate the reason for the anomaly, assess its impact on the quality of the observational data and determine possible implications for the ground segment. In severe cases, observations may be flagged as "failed" in the database, and made available for re-scheduling. For every observation, quality information will be gathered in a "quality control report summary" product, that is made available in the Herschel Science Archive, attached to the observational data. The report contains both the automatically generated quality control data and the conclusions of the problem analysis by the experts, when applicable. Items that will be included in the report are: MOC spacecraft and operations information, on-board observation execution anomalies (instrument or satellite related), telemetry gaps, pointing issues, space weather events, instrument specific warnings (e.g., high glitch rate), and data processing problems.

8.4. Herschel Science Archive

Herschel data products systematically generated by the SPG pipeline are made available to the users through the Herschel Science Archive (HSA) immediately after the pipeline processing is completed, typically 1-2 days after an observation has been executed. Following the completion of some basic quality control checks, something which may take from a few additional days to weeks, depending on the circumstances, notification e-mails are sent to the data owners that can then be informed about any quality issue affecting their observations.

By using the HSA User Interface, astronomers can search, browse, select and retrieve Herschel data products according to the observations proprietary rights as explained in Chapter 1 of the Herschel Data Analysis Guide.

Access to the HSA User Interface is provided from the following link:

http://archives.esac.esa.int/hsa/hsa.html

For every Herschel observation, the data populating the HSA consist of the observational products generated by the pipeline, containing the scientific data, together with the calibration and auxiliary

products as described in the Herschel Data Products section. In addition, associated quality information, generated to support archive users in the assessment of their scientific products, is also included.

The HSA provides data products as FITS format files at different levels of data reduction which can be used for further processing within the Herschel Interactive Processing Environment (HIPE), or with any other standard data processing package. It will also host highly processed data returned from the observers at a later stage in the mission.

8.5. Herschel Interactive Processing Environment

In addition to standard products, a software called Herschel Interactive Processing Environment (HIPE) is offered to the astronomical community to reduce the Herschel data interactively (starting from level-0, -1 or -2 products), and to perform science analysis on them.

The Herschel Interactive Processing Environment (HIPE) enables the user to:

• Access and retrieve data directly from the Herschel Science Archive, although it can also be retrieved independently from the HSA User Interface.

• Perform interactive data reduction from raw data to publishable products, using Herschel-provided and user-developed routines, both in GUI form or in console-batch mode. In particular, it contains the same pipeline scripts and tasks as the SPG so that users can reproduce the standard processing and add improvements to it where necessary.

• Visualise and manipulate image, spectral and spectral cube data.

• Perform science analysis with a number of built-in standard and configurable graphical and/or console-based tools.

• Get access to context-sensitive documentation and help.

The HIPE package does not require commercial licenses and is built to be platform-independent. It is based on Java and allows scripting programing in jython. The distribution includes source of software, calibration data and documentation. In addition, the astronomer will be able to develop and integrate his/her own data processing algorithms within the system.

Linux, Windows, or Mac installers for the latest user version of HIPE can be retrieved from the following link:

http://herschel.esac.esa.int/HIPE_download.shtml

Currently, a cycle of 4 releases per year is planned, to accommodate the fast evolution of the instrument knowledge and data-processing algorithms in the early phases of the mission.

HIPE is open to external contributions. HIPE pipelines are organised in modules (called tasks), easily interchangeable by user-customised tasks. The Key Programme consortia and the astronomical community in general are encouraged to feed back their data products and share the tools and algorithms developed to produce them with the HSC for possible inclusion in the Data Processing system.

Note
Data processing offers a number of interest groups for HIPE users. Details can be obtained from the Data Processing pages of the HSC Web page.

Chapter 9. Acronyms

2MASS - 2 Micron All-Sky Survey

AAS - Altitude Anomally Sensors

ACA - Altitude Control Axis

ACC - Attitude Control Computer

ACMS - Attitude Control and Measurement System

AGN - Active Galactic Nucleus

AME - Absolute Measurement Error

AOR - Astronomical Observation Request

AOT - Astronomical Observing Template

APE - Absolute Pointing Error

CP - Calibration Pointing

CFIRB - Cosmic Far Infrared Background

CRS - Coarse Rate Sensors

CUS - Common Uplink System

CVV - Cryostat Vacuum Vessel

DTCP - Daily Telecommunications Period

DSS - Deep Sky Survey

EPLM - Extended Payload Module

ESA - European Space Agency

ESAC - European Space Astronomy Centre

ESD - Electrostatic Discharge

ESOC - European Space Operations Centre

FIRSB - Far Infra Red Sky Background

FIRST - Far Infra Red Space Telescope

FoV - Field of View

FIR - Far Infra Red

FPU - Focal Plane Unit

FWHM - Full Width Half Maximum

GO - Geostationary Orbit

GYR - Gyroscope

HCNE - Herschel Confusion Noise Estimator

HIFI - Heterodyne Instrument for the Far Infrared

HOB - Herschel Optical Bench

HOTAC - Herschel Observing Time Allocation Committee

HSC - Herschel Science Centre

HIPE - Herschel Interactive Processing Environment

HST - Hubble Space Telescope

IA - Interactive Analysis

ICC - Instrument Control Centre

ID - Identification

IPAC - Infrared Processing and Analysis Center

IRAS - Infrared Astronomical Satellite

ISM - Interstellar Medium

ISO - Infrared Space Observatory

LEO - Low Earth Orbit

LEOP - Low Earth Orbit Phase

MOC - Mission Operations Centre

MIR - Mid InfraRed

MLI - Multi-Layer Insulation

NAIF - Navigation Ancillary Information Facility

NASA - National Aeronautics and Space Administration

NED - NASA Extragalactic Database

NHSC - NASA Herschel Sciencce Centre

OD - Operational Day

PACS - Photodetector Array Camera and Spectrometer

PDE - Pointing Drift Error

PDF - Portable Document Format

PLM - Payload Module

PSF - Point-source Spread Function

PV - Performance Verification

QCP - Quality Control Pipeline

RCS - Reaction Control System

RF - Radio Frequency

RPE - Relative Pointing Error

RWA - Reaction Wheel Assembly

S/C - Spacecraft

SAS - Sun Aquisition Sensors

SCUBA - Sub-millimetre Common-User Bolometer Array

SED - Spectral Energy Distribution

SPG - Software Product Generation

SPIRE - Spectral and Photometric Imaging REceiver

SREM - Standard Radiation Environment Monitor

SRPE - Spatial Relative Pointing Error

SSO - Solar System Object

SSR - Solid State Recorders

STR - Star Trackers

SVM - Service Module

TBD - To Be Determined

WFE - WaveFront Error

Chapter 10. Acknowledgements

This manual has been written by Pedro García-Lario, Bruno Merin, Miguel Sánchez-Portal and Mark Kidger (who acted as overall editor) and based on a structure initially produced by Timo Prusti, with additional inputs from Eva Verdugo and Asier Abreu.

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Chapter 11. Change record

- 2007/02/14: slight change Section 6.4.4 to clarify rules on permitted chainings.

- 2007/02/15: Inconsistency noticed between HSpot and PACS documentation. Short wavelength cut-off for PACS changed in Section 3.4 to be consistent with HSpot (also changed in PACS Manual by BA). Resolution information updated to give information on 1st, 2nd and 3rd order performance.

- 2007/04/11: Resolution information updated to give updated range information on 1st, 2nd and 3rd order performance in Section 3.4 to be consistent with values defined in SCR-3091.

- 2007/04/30: A sub-section is added Section 6.2.2.1 to explain the origin of NAIF IDs.

- 2007/05/29: Some typos corrected in Chapter 5.

- 2007/05/31:

  Updates of concatenation rules in Section 6.4.4.

  Updates of overhead rules and application in Section 6.6.

  Updates of calibration overhead rules and application in Section 6.6.3.

  Updates of constrained observation rules and application in Section 6.6.4.

  Correct HIFI exclusion half-angles in Table 2.3.

  Add SPIRE PACS parallel mode exclusion half-angles as a footnote in Table 2.3.

  Update PACS sensitivities in Table 3.2.

- 2007/08/01:

  Update to the proposal submission procedure in Section 6.1.3 to take into account the fact that proposers must now use the HerschelFORM PDFLatex pakage to prepare their scientific case.

  Update to the observing modes described in Section 6.2 to eliminate the cluster and shadow observing target types that currently seem unlikely to be implemented.

  Add a section Section 6.2.2.2 on the accuracy of the available ephemerides for moving targets.

  Add a section Section 6.2.2.3 on the required accuracy of ephemerides for moving targets for them to be observable by Herschel.

- 2010/04/30:

  Completed revised and updated for OT1 Call.

- 2010/07/08:

  Some small changes to Chapter 1.

- 2010/10/29:

  Added some notes on Solar System Object scheduling and ephemerid accuracy to Chapter 6.

- 2011/02/03:

  Added notes on response time for urgent schedule change requests and ToOs.

- 2011/03/31:

Complete revision and update of content in line with current knowledge.